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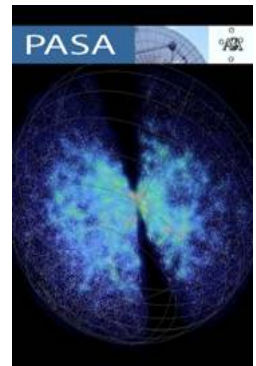
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Publications of the Astronomical Society of Australia / Volume 30 / January 2013 / e014  
DOI: 10.1017/pasa.2012.014, Published online: 23 January 2013

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### How to cite this article:

M. Todd, D. M. Coward, P. Tanga and W. Thuillot (2013). Australian Participation in the *Gaia* Follow-up Network for Solar System Objects. *Publications of the Astronomical Society of Australia*, 30, e014 doi:10.1017/pasa.2012.014

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# Australian Participation in the *Gaia* Follow-up Network for Solar System Objects

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(RECEIVED June 28, 2012; ACCEPTED August 29, 2012; ONLINE PUBLICATION January 24, 2013)

## Abstract

The *Gaia* satellite, planned for launch by the European Space Agency (ESA) in 2013, is the next-generation astrometry mission following *Hipparcos*. *Gaia*'s primary science goal is to determine the kinematics, chemical structure, and evolution of the Milky Way Galaxy. In addition to this core science goal, the *Gaia* space mission is expected to discover thousands of Solar System objects. Because of orbital constraints, *Gaia* will only have a limited opportunity for astrometric follow-up of these discoveries. In 2010, the *Gaia* Data Processing and Analysis Consortium (DPAC) initiated a program to identify ground-based optical telescopes for a *Gaia* follow-up network for Solar System Objects to perform the following critical tasks: confirmation of discovery, identification of body, object tracking to constrain orbits. To date, this network comprises 37 observing sites (representing 53 instruments). The Zadko Telescope, located in Western Australia, was highlighted as an important network node because of its southern location, longitude, and automated scheduling system. We describe the first follow-up tests using the fast moving Potentially Hazardous Asteroid 2005 YU55 as the target.

**Keywords:** instrumentation: miscellaneous – minor planets, asteroids: general – telescopes

## 1 INTRODUCTION

The European Space Agency (ESA) *Gaia* mission, scheduled for launch in spring 2013, is a space-based all-sky survey. The *Gaia* spacecraft will provide astrometry, photometry, and spectroscopy for point-like sources down to  $V \sim 20$ . *Gaia*'s science data comprise absolute astrometry, broad-band photometry, and low-resolution spectro-photometry. During its 5-year mission, *Gaia* will survey about 1 billion stars and 300 000 Solar System objects (SSOs), of which the majority will be main-belt asteroids. It will also survey about 500 000 point-like extragalactic sources and  $\sim 1$  million faint galaxies. The astrometric precision for the mission will be better than  $10 \mu\text{as}$  for stars brighter than  $V \sim 13$  and about  $25 \mu\text{as}$  for stars  $V \sim 15$ . *Gaia* will initially build on, and then surpass, the results of the *Hipparcos* mission of about 20 years ago (Mignard 2011). Such an observational effort has been compared to mapping the human genome, for the amount of collected data and for the impact that it will have on all branches of astronomy and astrophysics.

In addition to the above core science goal, the *Gaia* space mission will discover exotic transient objects in large numbers. Many thousands of transients will be discovered includ-

ing exoplanets and supernovae. Tens of thousands of brown and white dwarfs will be identified spectroscopically and, within our Solar System, some hundreds of thousands of minor planets will be observed. Of particular interest will be the numbers of unusual minor planets such as minor planets which have high inclinations, so that they are normally outside the regions of the sky routinely surveyed by near-Earth asteroid (NEA) program and inner Solar System Trojan asteroids.

Because *Gaia*'s primary mission is to perform a space-based all-sky survey, it is not designed to conduct any targeted follow-up studies. *Gaia* will be constrained by its orbit and by its design to survey the sky as completely as possible (Mignard et al. 2007). For these reasons, and because of the expected vast number of discoveries of transient phenomena, *Gaia* will not be able to either confirm or perform detailed studies of the discoveries.

### 1.1 *Gaia* Follow-up Network

During its 5-year mission, *Gaia* will observe many new SSOs. One feature of *Gaia* observations is the ability to image



**Figure 1.** The global distribution of observing stations participating in the *Gaia* FUN-SSO. The distribution is biased towards northern latitudes and European longitudes. Blue markers are the current observing stations, red markers are existing facilities which have expressed interest but are not confirmed, and yellow markers are planned future facilities.

at rather low solar elongation ( $45^\circ$ ), enabling detections of Earth-crossing asteroids (Atens) and inner-Earth asteroids (IEAs), and discoveries of new NEAs at larger solar elongation. In performing an all-sky survey, *Gaia* will necessarily survey regions of the sky away from the regions targeted by NEA programs. This holds the potential to discover asteroids with high orbital inclinations and Trojan asteroids in the stable Lagrangian regions in the orbits of the planets in the inner Solar System (Todd et al. 2011, 2012a, 2012b).

In such cases, due to the motion of the objects and limiting magnitude, the scanning law of *Gaia* will restrict the orbit determination to be constrained on a very small number of astrometric measurements. This implies that only a ground-based optical network can obtain accurate orbital modelling, based on enough astrometric measurements. This is the primary objective of the *Gaia* Follow-up Network for Solar System Objects (FUN-SSO).

In addition to the improvement in accuracy of the astrometric data used for orbital modelling of specific objects, some rare and peculiar SSOs such as asteroids with cometary activity could be studied (Tanga et al. 2008). Due to the limitations of *Gaia*'s observing method, a ground-based follow-up network will be crucial for studying the physical characteristics of these objects.

In order to be effective in acquiring high-quality follow-up imaging following an alert by the *Gaia* data processing system, the network must have a large geographical coverage. This is why several observing stations have been in-

vited for participation in this project. In 2010, the *Gaia* Data Processing and Analysis Consortium (DPAC) initiated a programme to identify optical telescopes for a *Gaia* follow-up network for SSOs to perform the following critical tasks: confirmation of alert, identification of transient, continuous monitoring, and tracking. At the *Gaia* Follow-up Network Meeting in Paris 2010, the Zadko Telescope was highlighted as an important contribution to this network, because of its longitude and robotic operation. To date, the network comprises 37 observing sites (representing 53 instruments) (cf. Figure 1).

SSOs and stars are not differentiated on board but are treated the same way. Each source brighter than  $V = 20$  is identified by the star mapper on the CCD array. A window is defined for each identified source for accumulation of charge as it crosses the CCD. An automated data reduction pipeline (implemented by the science consortium, DPAC) is run over all sources. Those which are found to move (from one epoch to the other, essentially) or are extended are sent to a specific branch of the pipeline which applies a specialised data reduction including identification and processing of SSOs. Most of the SSOs that *Gaia* will observe will be known, so they will be simply identified by comparison to ephemeris derived from the complete Minor Planet Center (MPC) catalogue. The remaining objects which are moving and not identified are classed as new asteroid candidates. For these, a tentative orbit is computed for ground-based recovery and an alert will be issued to the *Gaia*-FUN-SSO network.

Alerts from the *Gaia* data processing system will be received between 24 and 48 h after detection. Problems may arise for observations of peculiar objects: fast moving objects, faint objects, and NEAs close to Earth with strong parallaxes. Hence, small and less sensitive telescopes (<0.6 m) can be useful but more sensitive instruments (>1 m) will be essential. Furthermore, fully robotic telescopes are proved as ideal for observations triggered by an alert. Among the telescopes in the *Gaia*-FUN-SSO network, five are robotic telescopes.<sup>1</sup>

## 2 THE ZADKO TELESCOPE

The Zadko Telescope<sup>2</sup> is a 1-m f/4 Cassegrain telescope, built by DFM Engineering, Inc.,<sup>3</sup> and is situated in the state of Western Australia at longitude 115°42'49'' E, latitude 31°21'24'' S, and at an altitude of 47 m above sea level (67 km north of Perth).

It is an equatorial, fork-mounted telescope with a primary mirror clear aperture of 1 007.0 mm, and a system focal length of 4 038.6 mm. Its fast optics have a low f-ratio, coupled with a flat, wide field of view (maximum possible FoV  $\approx 1^\circ$ ).

The core science theme of the Zadko Telescope is the discovery and study of transients (Coward et al. 2010). The facility was robotised in 2010, employing an observatory control system developed for the *Télescopes à Action Rapide pour les Objets Transitoires* (TAROT). It can acquire photometric data autonomously on gamma-ray bursts (GRBs) and Galactic transients following alerts by the *Swift* satellite and other external triggers. In addition to the follow-up of GRB triggers, the Zadko Telescope is used for the follow-up of exoplanet discoveries, follow-up imaging of Astronomy with a Neutrino Telescope and Abyss environmental RE-Search (ANTARES) alerts, follow-up imaging of Laser Interferometer Gravitational Wave Observatory (LIGO)/Virgo gravitational wave alerts, and SSOs. Some Zadko Telescope observations are as follows.

- Zadko measures the light curve of GRB 101024 only 200 s after satellite trigger, revealing unexplained optical light-curve features (Gendre et al. 2011).
- In 2009, a minor planet pilot search programme resulted in the discovery of 13 uncatalogued SSOs. Further studies resulted in four of these being linked to ‘orphan’ observations in the MPC database.
- Fourteen follow-ups of *Swift* satellite GRB triggers with NASA Gamma-ray Coordinates Network notices.
- The TAROT + Zadko network successfully images sky positions triggered by LIGO for the detection of possible optical counterparts to gravitational wave candidates (LIGO Scientific Collaboration et al. 2012).
- GRBs 090205 and 090516, with respective redshifts of 4.3 and 4.1, are among the most distant optical transients imaged by an Australian telescope.

<sup>1</sup> See <http://www.imcce.fr/gaia-fun-sso/> for an updated list.

<sup>2</sup> <http://www.zt.science.uwa.edu.au>

<sup>3</sup> <http://www.dfmengineering.com>

**Table 1.** Zadko Telescope Asteroid Discovery and Follow-up Statistics

Year	Detections	No. of Objects Linkages	Link Span (opp. s/years)
2009	13	4	4/10
2010	4	1	3/5
2011	6	1	5/5
2012	1	1	2/6

For *Gaia* to achieve astrometric accuracy on scales of a few tens of microarcseconds, its reference frame must be well defined. By using the optical positions of quasi-stellar objects (QSOs), assuming that QSOs have no peculiar transverse motion, then any observed proper motion will reveal the global rotation of the *Gaia* sphere. Along with the robotic TAROT telescopes in France and Chile, the Zadko Telescope has been observing selected QSOs since mid-2011 to assist in constructing the initial catalogue for establishing *Gaia*'s reference frame.

The Zadko Telescope observatory is being upgraded in 2012. The observatory building is being replaced with a new observatory optimised for automated rapid response imaging of transients and SSOs. Furthermore, a new instrument and imaging package with spectroscopic capability, funded by the Australian Research Council, will be installed in 2012 and 2013.

The capability of the Zadko Telescope to detect faint objects to  $m \approx 21.5$  and its robotic scheduling system enables it to perform optimal targeted follow-up observations. Since 2010, routine examination of selected image sets where the cadence has facilitated detection of moving objects has resulted in the detection of 11 additional uncatalogued SSOs. Follow-up studies have enabled three of these to be linked to earlier detections in the MPC database. In some cases, these linkages have been made between apparitions at several previous oppositions over a period of years up to a decade (see Table 1).

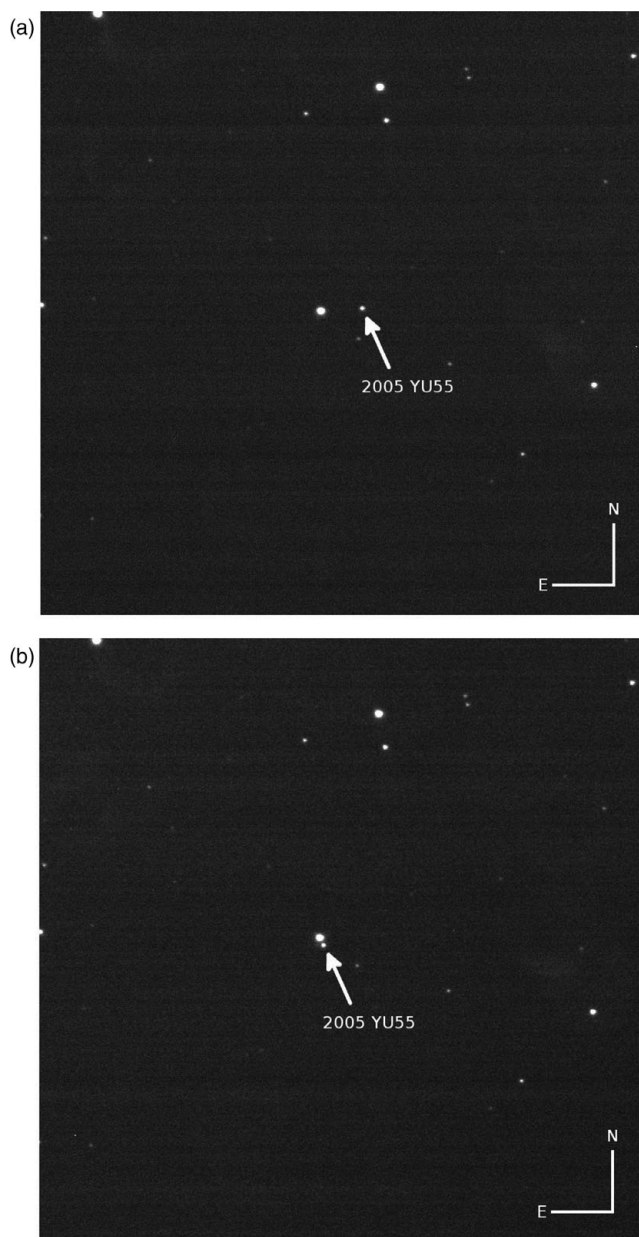
## 3 PRELIMINARY *Gaia* FUN-SSO TESTING

In 2011 November, observations were made of the Potentially Hazardous Asteroid (PHA) 2005 YU55. On 8 November, this object passed within 325 000 km of Earth.<sup>4</sup> Observing this target so soon after its fly-by of Earth meant that its position was rapidly changing (see Figure 2). An observation of this target demonstrated the capability of the Zadko Telescope to accurately point and acquire images of a relatively fast-moving object. On the first night of the observation, the sky motion was  $\sim 4$  arcsec min<sup>-1</sup>. By comparison, main-belt asteroid sky motions are typically smaller than 0.5 arcsec min<sup>-1</sup>.

The Zadko Telescope made observations of 2005 YU55 over a period spanning 11 nights, from 11 to 21 November.

<sup>4</sup> [www.jpl.nasa.gov/news/news.cfm?release=2011-332](http://www.jpl.nasa.gov/news/news.cfm?release=2011-332)





**Figure 2.** Zadko Telescope images of PHA 2005 YU55. Exp: 60 s, field centre: 02h20m49s6 +17d06m47s, scale: 11.5 × 11.5 arcmin, date: 2011-11-12.60964 (a) and 2011-11-12.62607 (b).

This was the third-longest observational coverage of 14 participating telescopes, and was the only contribution from a Southern Hemisphere telescope. The Zadko Telescope was used to obtain 42 images of PHA 2005 YU55 in 2011 November, of which 21 were used for astrometry. The images and measurements were delivered to the *Gaia* FUN-SSO working group, which reported the collective results from participating instruments to the MPC. During this time, the target faded dramatically (from 13.7 to 18.1 in the Sloan Digital Sky Survey (SDSS) *r* filter).

The residual error between the reported (observed) position and the position calculated by the *Gaia* FUN-SSO

working group was calculated by taking ~3 600 observations between 2005 December and 2011 December and fitting a numerical model of the asteroid motion. The model takes into account perturbations from all eight major planets and Pluto and the relativistic effect of the Sun. The mean residual was determined to be 0.30 arcsec, and for the Zadko Telescope was typically better than 0.20 arcsec in both right ascension and declination. The first measurement from each night, with residuals ( $O - C$  values), is given in Table 2, and shows the change in brightness and sky motion over that period.

In 2012 February, observations were planned for asteroids 99942 Apophis and 1996 FG3. Apophis was selected since it will have a very close approach in 2029 and new astrometry would be useful to obtain more accurate ephemerides. New data can also be used to improve our knowledge of the Yarkovsky effect on its trajectory. 1996 FG3 was selected as it is the main target of the proposed Marco Polo-R space mission which will provide a sample return. Better knowledge of its orbital characteristics will help with the success of the mission. Although no data were obtained for these because of observing constraints, the targets were successfully logged into the robotic scheduler as a test.

#### 4 SUMMARY

During its 5-year mission, *Gaia* will discover new SSOs, including IEAs and Earth-crossers (Atens), and new NEAs at larger solar elongations. Because the scanning law of *Gaia* will restrict the orbit determination to a very small number of astrometric measurements, only a ground-based optical network can obtain accurate orbital modelling, based on enough astrometric measurements. This is the primary objective of the *Gaia* FUN-SSO.

In 2010, the *Gaia* DPAC initiated a programme to identify ground-based optical telescopes for the *Gaia* FUN-SSO. The main criteria for the network are facilities that are ideally robotic, or can be re-scheduled to give priority to a *Gaia* alert. The Zadko Telescope was identified as a potentially important contributor to the network because of its robotic operation, sensitivity, and location. It will be used for validation of new SSOs, astrometry, and photometry of these targets.

In 2011, the *Gaia* FUN-SSO working group initiated an observing campaign of SSOs, with the main goal to test the capability of the network node facilities. Simulated *Gaia* alerts were distributed to the network (initially by email) providing ephemerides for the follow-up imaging of real targets. We tested the Zadko Telescope capability for participating in the observing campaign of PHA 2005 YU55 in late 2011. Of the 37 observing stations in the network, 11 were able to participate and provide useful astrometric data. The Zadko Telescope contributed the third-longest observational coverage from 14 participating telescopes and was the only contribution from a Southern Hemisphere telescope. The data from the observations showed that the Zadko Telescope scheduling system worked successfully, with accurate telescope

**Table 2.** Selection of 2005 YU55 Observations Made by the Zadko Telescope Showing the Change in Brightness and Sky Motion Over the Period. The  $O - C$  Values Are the Differences from the Calculated Positions

Date (UT)	RA	Decl.	Mag.	Sky Motion (arcsec min <sup>-1</sup> )		$O - C$ (arcsec)	
				$\alpha$	$\delta$	$\alpha\cos(\delta)$	$\delta$
2011 11 11.60593	02 13 56.45	+17 24 22.7	13.7	+4.1	-0.9	-0.43	-0.10
2011 11 12.60797	02 20 47.59	+17 06 54.4	14.2	+1.9	-0.5	-0.20	-0.09
2011 11 15.60574	02 29 03.05	+16 45 08.8	16.1	+0.3	-0.1	0.05	-0.12
2011 11 16.60259	02 30 24.47	+16 41 42.2	16.3	+0.2	-0.09	0.01	0.04
2011 11 19.59546	02 33 06.84	+16 35 40.7	17.7	+0.03	-0.03	-0.00	-0.13
2011 11 21.59059	02 34 22.04	+16 33 47.7	18.3	+0.02	-0.02	0.50	-0.02

pointings, and positional data that were used for astrometry with an uncertainty of about 0.2 arcsec. Testing of the networks ability to constrain orbits of SSOs will continue until the launch of *Gaia*.

Following upgrades to the Zadko Telescope observatory in 2012 and 2013, the telescope will resume operation for several optical transient projects. These improvements will also provide scope for participation in additional projects, including the follow-up of exoplanet candidates. Furthermore, an image processing pipeline for automated identification and analysis of optical transients is in development and planned for implementation in 2013.

## ACKNOWLEDGMENTS

D. M. Coward is supported by an Australian Research Council Future Fellowship. The Zadko Telescope was made possible by a philanthropic donation to the University of Western Australia by

the Zadko family. M. Todd and D. Coward acknowledge support received from the European Science Foundation (ESF) for the activity entitled 'Gaia Research for European Astronomy Training'.

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