

We associate the northern point source with K308, which was identified by Kaas et al. (2004) as a Class I object. It is not included in HMH07 due to the lack of a detection at 24 μm . Conversely, Winston et al. (2007) identify this object as Class II; however here we retain the Class I designation due to the presence of a sub-mm core and we identify this source with Ser-emb 22 (Enoch et al. 2009a). The extended emission to the south of DCE08-215 is more difficult to divide conclusively. We tentatively divide this emission between the two Class I sources K331 (Kaas et al. 2004) associated with VLA-19 (Eiroa et al. 2005) and HMH07-198, and K330 (Kaas et al. 2004) associated with HMH07-197.

4.2 Cluster B Region

4.2.1 Existing radio observations of Serpens Cluster B

Observations of the Serpens Cluster B region were made with the VLA at 3.5 cm by Djupvik et al. (2006). Of nine radio detections we associate four with protostellar objects from HMH07. As with Cluster A the majority (three) of these radio sources correspond to Class 0/I objects; the remaining detection is associated with HMH07-59, which is identified as Class II/III (HMH07; Harvey & Dunham 2009).

4.2.2 DCE08-208

The object DCE08-208 was identified as an external galaxy candidate by EDJ09 ([EDJ2009]-577) with a bolometric luminosity of $0.024 L_{\odot}$, but is considered a Group 1 candidate for an embedded protostellar source by DCE08. HMH07 who initially identify this object as HMH07-75, divide it further into three components A, B and C (Harvey et al. 2007b). Of these components they make preliminary classifications based on bolometric temperature of Class

I, I and 0, respectively. Component C has a 70 μm counterpart, whereas components A and B do not, and it has been suggested that, although component A is point-like, component B may simply be extended emission from the outflow of component C. The combined bolometric luminosity of the three sources is heavily dominated by the Class 0 component, C, which has a re-derived value of $1.6 L_{\odot}$ (Harvey & Dunham 2009).

Radio emission from DCE08-208 is detected at 16 GHz as an extension to a brighter adjacent source (see Fig. 4a). It is likely that this emission is dominated by HMH07-75C, and appears to be extended towards the north of this source. There is no corresponding radio source detected at 3.5 cm (Djupvik et al. 2006). Assuming that DCE08-208 was not significantly removed from the phase centre of the 3.5 cm VLA observations this would give the radio emission a spectral index of $\alpha \geq 1.13$.

Immediately to the south-west of DCE08-208 is a bright radio source coincident with objects HMH07-67 and HMH07-68. We associate this source with the Class 0 object HMH07-68 (Harvey & Dunham 2009), based on the proximity of the radio peak to the *Spitzer* position; however it may also have a contribution from the closely removed Class I YSO HMH07-67. HMH07-68 was also detected at 3.5 cm by Djupvik et al. (2006) and using their flux density we measure a spectral index for this object of $\alpha = 2.07 \pm 0.40$. Such a steep rising spectral index is typical of optically thick free-free emission, such as is seen towards compact H II regions. These types of massive H II regions are not typical for Class 0 protostellar objects and it is therefore more likely that the significant increase in flux between 3.5 and 1.8 cm is due to the presence of either extended structure around this object which has been lost in the higher resolution VLA observations, or alternatively variability in the radio emission from HMH07-068.

Further to the south-west there is a third radio source, which is centred on the two sources HMH07-58 and HMH07-59. HMH07-59

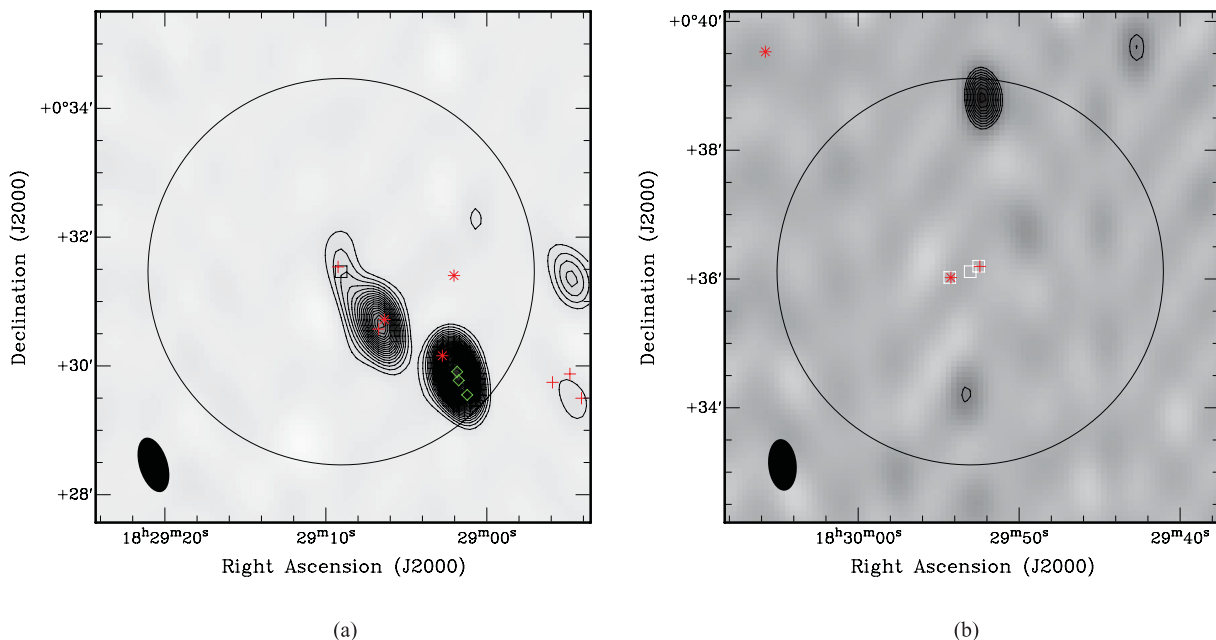


Figure 4. Serpens Cluster B fields. (a) DCE08-208 field; (b) DCE08-213 field (no protostellar sources detected). Grey-scale and contours are from the AMI-LA at 16 GHz, with contours at 5, 6, 7, 8 σ etc. where $\sigma = 16 \mu\text{Jy beam}^{-1}$ for (a) and $\sigma = 19 \mu\text{Jy beam}^{-1}$ for (b). Crosses ('+') mark the positions of Class 0 protostellar objects, stars ('*') mark the positions of Class I objects, and diamonds ('◇') mark the positions of Class II sources (see Table 5). The position of candidate embedded objects from DCE08 is shown as squares. The AMI-LA primary beam FWHM is shown as a solid circle and the synthesized beam as a filled ellipse in the bottom-left corner.

was also detected at 3.5 cm by Djupvik et al. (2006) and it was suggested that the emission from this more evolved Class II/III protostellar object could be due to gyro-synchrotron emission. If we assume that the flux density seen at 16 GHz is entirely due to HMH07-59 we find that this object has a slightly rising spectrum with $\alpha = 0.13 \pm 0.04$. Again such a spectrum could imply free-free as the dominant emission mechanism; however it is also not inconsistent with the previously proposed emission mechanism of gyrosynchrotron and this is discussed further in Section 6.4. The elongated nature of the detection suggests that it may also have contributions from HMH07-54 and HMH07-56, which are again identified as Class II/III objects. Fitting the 16 GHz emission as the sum of two point sources, the first centred on HMH07-59 and the second on HMH07-56, we can constrain this additional emission to be ≤ 0.22 mJy. The fact that this second component was not detected at 3.5 cm implies that it also possesses a rising spectrum.

4.2.3 DCE08-213

There are two Class 0 objects close to the pointing centre of the DCE08-213 AMI-LA field. Neither of these is detected, and indeed no protostellar sources are detected at 16 GHz in this field with flux densities $> 96 \mu\text{Jy}$. Two radio point sources are detected away from the pointing centre. The radio source to the north is identified with NVSS J182952+003857 and has $S_{1.4\text{GHz}} = 2.8 \pm 0.6$ mJy. Combining these data with the AMI-LA 16 GHz flux density shows that this is a steep spectrum extragalactic radio source with $\alpha = -0.58 \pm 0.07$. The second source has a flux density approximately half that of the northern source and, assuming a similar spectral index, this would place it below the detection limit in the NVSS survey at 1.4 GHz. With no known protostellar object in the vicinity of this source we identify it as a non-thermal extragalactic radio source.

4.3 Serpens filament

4.3.1 DCE08-202

The DCE08-202 field in the Serpens filament is completely devoid of radio sources at 16 GHz (see Fig. 5). The nearest source is almost 4 arcmin to the east of the pointing field and is associated with NVSS J182901+005117. Flux densities measured this far outside the AMI-LA primary beam are imprecise due to uncertainties in the primary beam model outside the FWHM but this source appears to have a flat spectrum $S_{1.4\text{GHz}} = 3.8 \pm 0.5$ mJy, $S_{16\text{GHz}} = 3.5 \pm 0.4$ mJy assuming 10 per cent calibration uncertainty at 16 GHz gives $\alpha = -0.03 \pm 0.01$.

The three protostellar objects in this field are all Class I and are uniformly not detected. In the case of HMH07-26 this is surprising as its bolometric luminosity would suggest that it should be detectable according to the correlation between radio and bolometric luminosity measured in Paper II. However, it is immediately noticeable from Table 6 that although these objects vary significantly in bolometric luminosity they all have similarly low envelope mass estimates. In the case of HMH07-26, which when judged on its bolometric luminosity would have a 16 GHz flux density of almost 300 μJy , has only a predicted 16 GHz flux density of 38 μJy when judged on its envelope mass.

4.4 Expected contamination by extragalactic radio sources

At 16 GHz we expect a certain number of extragalactic radio sources to be seen within each of our fields. Following Papers I and II,

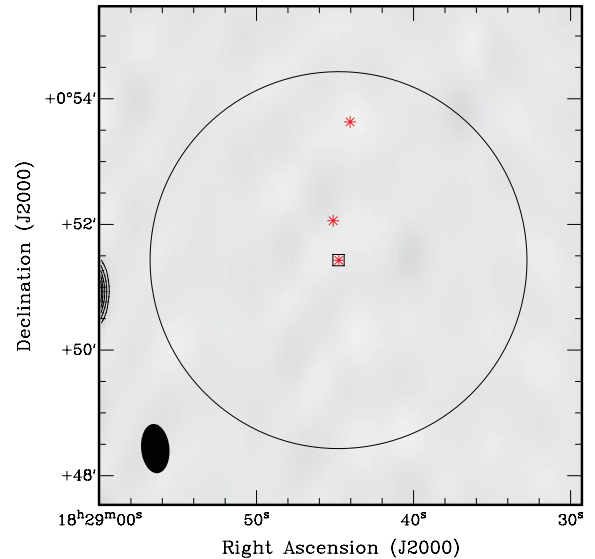


Figure 5. DCE08-202 (no protostellar sources detected). Grey-scale and contours are from the AMI-LA at 16 GHz, with contours at 5, 6, 7, 8 σ etc. where $\sigma = 20 \mu\text{Jy beam}^{-1}$. Stars ('*') mark the positions of Class I objects (see Table 6). The position of the candidate embedded object from DCE08 is shown as a square. No Class 0 objects are present in this field. The AMI-LA primary beam FWHM is shown as a solid circle and the synthesized beam as a filled ellipse in the bottom-left corner.

to quantify this number we use the 15 GHz source counts model from de Zotti et al. (2005) scaled to the 10C survey source counts (Davies et al. 2010). The average rms noise from our data sets is $\approx 19 \mu\text{Jy beam}^{-1}$ and from this model we predict that we should see 0.07 sources arcmin^{-2} , or ≈ 2 radio sources within a 6 arcmin FWHM primary beam above a 5σ flux density of 95 μJy . Within the Serpens core region we would therefore expect to detect $\approx 4 \pm 2$ extragalactic radio sources, and in the Serpens/G3-G6 region a further $\approx 9 \pm 3$ sources. Making the assumption that all sources which cannot be identified with a previously known protostellar object are extragalactic, as are sources with non-thermal spectra, we find three radio sources in the Serpens Cluster A fields, consistent with our prediction; in the remaining fields we find three further sources, which is lower than predicted, but consistent at 2σ . The low number of extragalactic sources detected in these fields compared to the model may represent an overestimation from the 15 GHz source counts, which are extrapolated from a completeness limit of 0.5 mJy – high compared with our detection limit of 95 μJy . One break in the observed source counts is already known at mJy flux density levels (Davies et al. 2010) and it is quite possible that a further break exists below 0.5 mJy.

4.5 Expected radio flux density from thermal dust emission

At 16 GHz there is expected to be a small contribution to the radio flux density of protostars due to the long wavelength tail of the thermal dust emission from the envelopes around these young objects, and this contribution may be estimated using sub-mm data. However, the complexity and multiplicity of objects in Serpens makes accurate flux density calculation from sub-mm maps difficult; sources are crowded and several sources with near neighbours may be associated with the same sub-mm core. For the purposes of this work we therefore constrain the contribution of thermal dust emission to the measured 1.8 cm flux densities of our detected

sources by extrapolating modified greybody spectra for each object directly from 1.1 mm flux densities. As the flux densities themselves are not listed by Enoch et al. (2009a) we recover them from the envelope masses in that work for Class 0 and Class I objects in Serpens, following

$$S_{1.1\text{ mm}} = \frac{M_{\text{env}} B_{1.1\text{ mm}}(T_d) \kappa_{1.1\text{ mm}}}{D^2}, \quad (3)$$

where $D = 260\text{ pc}$ is the distance to Serpens, $\kappa_{1.1\text{ mm}} = 0.0114\text{ cm}^2\text{ g}^{-1}$ is the opacity and $B_{1.1\text{ mm}}(T_d)$ is the Planck function at a dust temperature of $T_d = 15\text{ K}$. The values of these parameters are chosen to match those used by Enoch et al. (2009a) to derive envelope mass estimates from flux densities. The opacity is taken from the dust model of Ossenkopf & Henning (1994) for grains which have coagulated over 10^5 years and which have thin ice mantles (OH5; indicating column 5 of table 1 in Ossenkopf & Henning 1994).

The source HMH07-139 (AMI-4) has no counterpart in the Bolocam catalogue of Enoch et al. (2009a); however it is coincident with a distinct sub-mm peak in the higher resolution 1.3 mm MAMBO map of Kaas et al. (2004) where it is identified as a Class I object. Following the method of Enoch et al. (2009a) we extract a sub-mm flux density for HMH07-139 by integrating the flux density within a 30 arcsec aperture, centred on the *Spitzer* position, from the publically available $850\text{ }\mu\text{m}$ data of Di Francesco et al. (2008). From these data we measure a flux density, $S_{850\text{ }\mu\text{m}} = 1.9\text{ Jy}$, and consequently an envelope mass of $M_{\text{env}} = 1.16\text{ M}_{\odot}$ (see Table 4).

Consequently, only one of the Class 0/I radio detections lacks a sub-mm flux density and envelope mass estimate, HMH07-186. This source is not identified with a distinct sub-mm peak, but instead is located between several adjacent compact cores. Without evidence for a peak at the position of this source we do not attempt to fit a flux density and envelope mass.

From the 1.1 mm flux densities we estimate the thermal dust contribution at 16 GHz by extrapolating a modified greybody model of the form $S_{\nu} = \nu^{\beta} B_{\nu}(T_d)$ from the 1.1 mm measurement. To remain consistent with the envelope mass estimates we again use a dust temperature of $T_d = 15\text{ K}$ and an index of $\beta_{\text{OH5}} = 1.85$. Using an alternative dust temperature of 10 K would increase the flux densities by approximately 30 per cent. The predictions themselves have uncertainties of 20–30 per cent from the errors on the 1.1 mm flux densities alone. We note also that, as indicated in Ossenkopf & Henning (1994), it is unlikely that β will remain constant from

1.1 mm to 1.8 cm. However in the absence of additional data it is not currently possible to constrain this model further.

In what follows we use the measured flux density at 1.8 cm with the predicted greybody contribution *subtracted* to calculate the radio luminosity. This is to ensure that the values used are representative only of the radio emission and do not include contributions from the thermal dust tail which varies greatly between sources and which might therefore influence any conclusions being drawn from the correlations examined in the later stages of this paper. In general, the corrections made to the radio flux densities for this sample are small, in the range 1–5 per cent. Notable exceptions are AMI-3 and AMI-7 with corrections of 20 and 10 per cent, respectively. We assume that any thermal dust contribution at 3.5 cm is negligible and make no correction for this quantity in the flux densities taken from the literature.

5 CORRELATIONS

We consider both the 3.5 and 1.8 cm data for these fields when forming correlations with the IR and sub-mm characteristics of the detected objects. In doing this we make two adjustments to the measured data. First, for the 3.5 cm flux densities we apply a scaling factor of $1.49 = (3.5/1.8)^{0.6}$, in order to compare them directly with the 1.8 cm flux densities. A spectral index of $\alpha = 0.6$ corresponds to partially optically thick free-free emission from both a spherically symmetric density distribution (Panagia & Felli 1975) and a collimated density outflow distribution (Reynolds 1986) and is generally assumed for the radio emission from protostellar objects. Secondly, we scale the envelope mass values from Hatchell et al. (2007) used in Paper II by a factor of 0.33, in order to compare them directly with the mass values for objects in Serpens which are taken from the Bolocam catalogue of Enoch et al. (2009b). This factor is the average ratio of envelope masses for the sources from the Perseus region studied in Paper II when measured with Bolocam (Enoch et al. 2009b) and with Submillimetre Common-User Bolometer Array (SCUBA; Hatchell et al. 2007). This factor accounts for differences in the assumed dust temperature, T_d , the opacity, κ_{ν} , after correcting for the assumed distance to Perseus, between these works when calculating mass values. We note that the limiting mass for the Enoch et al. (2009b) Bolocam catalogue is $M_{\text{env}} \simeq 0.1\text{ M}_{\odot}$, significantly below the lower limit indicated in Figs 6 and 7.

It can be seen (Fig. 6) that the 3.5 cm data are consistent with the correlations measured in Papers I and II. There are two notable

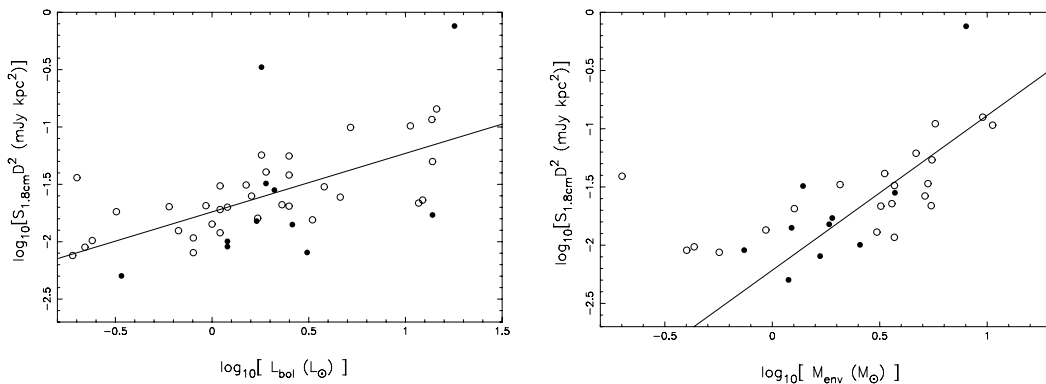


Figure 6. (a) Correlation of bolometric luminosity with 1.8 cm radio luminosity; (b) correlation of envelope mass with 1.8 cm radio luminosity. Unfilled circles show data from Papers I and II in (a) and from Paper II in (b). Filled circles show 3.5 cm data from Eiroa et al. and Djupvik et al. 3.5 cm radio flux densities have been scaled by a factor of 1.49 (see text for details); masses from SCUBA data (H07; Paper II) have been scaled by a factor of 0.33 to match those derived from Bolocam (E09; see text for details). Best-fitting correlations from Paper II are shown as solid lines.

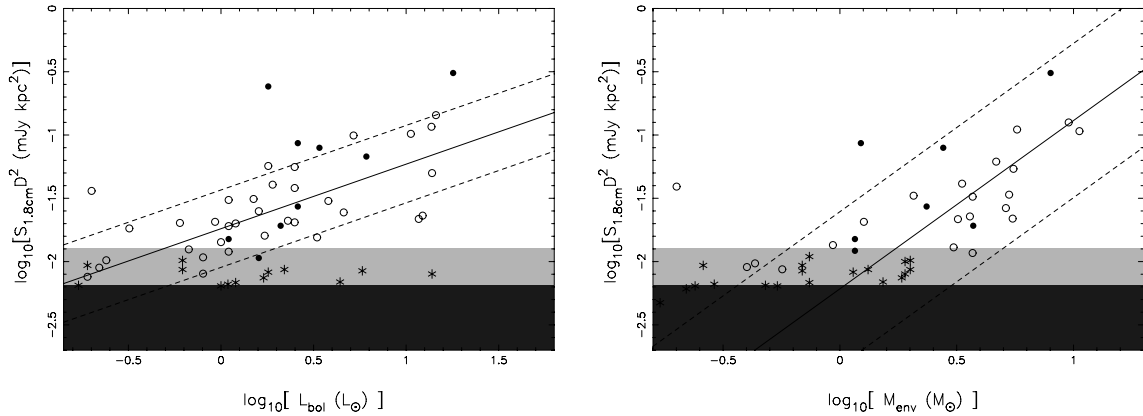


Figure 7. (a) Correlation of bolometric luminosity with 1.8 cm radio luminosity; (b) correlation of envelope mass with 1.8 cm radio luminosity, where masses from SCUBA data (H07; Paper II) have been scaled by a factor of 0.33 to match those derived from Bolocam (E09; see text for details). Data points from Papers I and II are shown as open circles; data points from this work are shown as filled circles for detected sources and stars (*) at the position at the 5σ upper limit for non-detections. Best-fitting correlations from Paper II are shown as solid lines, with the 1σ bounds on these fits indicated by dashed lines. Radio luminosities below the detection limit at the phase centre are shaded in dark grey, and the region shaded in light grey indicates the range of radio luminosities which lie between the detection limit at the phase centre and the detection limit at the half power point of the primary beam.

outliers in Fig. 6, DCE08-210 and HMM07-59. Bolometric luminosity estimates for DCE08-210 (Serpens SMM 1) vary; here we adopt a bolometric luminosity for this source of $L_{\text{bol}} = 17.3 L_{\odot}$ (Enoch et al. 2009b). This value is higher than that determined solely from broad-band sub-mm flux densities (Enoch et al. 2009a, see Table 4) as it is calculated using the radiative transfer model of Dullemond & Dominik (2004) and takes into account inclination effects.

DCE08-210 is also detected in the NVSS survey at 1.4 GHz (Condon et al. 1998) with a flux density $S_{1.4\text{GHz}} = 4.5 \pm 0.5$ mJy implying that it has a rising spectrum from 21 to 1.8 cm. The flux density measured at 3.5 cm is in excess of both the 21 and 1.8 cm values, preventing a single power law fit. This may imply that DCE08-210 was ‘flaring’ at the epoch of the 3.5 cm VLA observations, which could explain the deviation of this datum from the general correlation. HMM07-59 is identified in Harvey & Dunham (2009) as a Class II/III source and consequently is the only object of its type included in this sample. It is not certain therefore that such an object should follow the general trend observed for Class 0 and Class I objects.

6 DISCUSSION

6.1 Detections and non-detections

We show both the detected and undetected sources in Figs 7(a) and (b). In these plots data from Papers I and II are shown as unfilled circles, detections from this work are shown as filled circles, and non-detections from this work are shown as stars at the position of their 5σ detection limits as listed in Tables 4–6. The plot area below the average 5σ detection limit at the field centres for the full sample of $95 \mu\text{Jy}$ is shaded in dark grey. The area corresponding to the detection limit between the field centres and the primary beam FWHM is shaded in light grey. Using the distribution of the data from Papers I and II, 1σ bounds are shown as dashed lines for the measured correlations, which are shown as solid lines.

As has already been remarked upon, in Fig. 7(a) there are two detected sources which lie significantly above the bounds on the measured correlation. These are DCE08-210 (Serpens SMM 1) and HMM07-59. Both of these objects also have discrepant 3.5 cm flux

densities (see Fig. 6a). HMM07-59 is a Class II/III source and, as the only source of this kind in the data, it is uncertain whether HMM07-59 should follow the general trend derived for Class 0/I objects. Since the radio emission mechanism for Class II/III objects is unknown, and considered to be different from that of Class 0/I, it seems possible that this object could deviate from the correlation. The 3.5 cm flux density for DCE08-210 has already been suggested to represent a ‘flare’ in the emission of this object, a possibility that is supported by the increased removal of this source from the correlation at 3.5 cm compared with 1.8 cm. The bolometric luminosity used here is $L_{\text{bol}} = 17.3 L_{\odot}$; however significantly higher luminosities have been reported for this object which may explain the slight offset of this object from the trend.

In Fig. 7(a), there are a number of undetected sources which lie below the detection threshold, but outside the bounds on the general trend, in the region $L_{\text{bol}} > 2.0 L_{\odot}$. There are four undetected sources in this region; three of these are located in multiple fields and in this case the limit corresponding to the field where the source is closest to the phase centre is plotted. Noticeably, however, this population does not appear correspondingly discrepant when considered in terms of their envelope mass (Fig. 7b). A possible reason for this is the high ratio of bolometric luminosity to envelope mass measured for these objects. Excluding these four objects the mean ratio, $r = L_{\text{bol}}/M_{\text{env}}$, for this sample is $r = 1.13$. HMM07-182 and HMM07-176, which have the highest luminosities of those sources that are undetected, have $r = 7.26$ and $r = 8.41$, respectively. The remaining two objects in this region, HMM07-137 and HMM07-208, have $r = 2.87$ and $r = 1.67$, respectively.

Neglecting uncertainties in the assumed distance to Serpens, which are commented upon later, the errors on radio luminosity in these correlations are dominated by the conservative absolute calibration error of the AMI-LA of 5 per cent. Uncertainties on the envelope mass have contributions from the assumed opacity index and temperature, as well as a 20–30 per cent uncertainty from the sub-mm flux density, whereas errors on the bolometric luminosity, as stated in Enoch et al. (2009a), are approximately 10 per cent.

The relationship of L_{bol} and M_{env} is often used to describe the evolution of embedded YSOs. The $L_{\text{bol}}-M_{\text{env}}$ diagram for this sample is illustrated in Fig. 8. We make a simple illustration of the evolutionary behaviour of the YSOs in this diagram under

Table 5. AMI-LA observations of the Serpens Cloud B. Columns are (1) AMI-LA source number; (2) right ascension of radio source peak; (3) declination of radio source peak; (4) integrated 1.8 cm flux density; (5) VLA association; (6) integrated 3.5 cm flux density; (7) spectral index from 3.5 to 1.8 cm; (8) associations from HH07; (9) bolometric luminosity from Enoch et al. (2009a), values in italics from EDJ09; (10) envelope mass from Enoch et al. (2009a) and (11) protostellar classification.

AMI	RA (J2000)	Dec. (J2000)	$S_{1.8\text{cm,int}}$ (mJy)	VLA	$S_{3.5\text{cm,int}}$ (mJy)	$\alpha_{3.5}^{1.8}$	[HH07]	L_{bol} (L_{\odot})	M_{env} (M_{\odot})	Cl.
<i>DCE08-208:</i>										
9	18 29 00.8	+00 32 19	0.141 ± 0.023	–	–	–	–	–	–	–
10	18 29 01.7	+00 29 48	3.583 ± 0.062	4	3.30	0.13 ± 0.04	59	4.2^a	–	II/III
11	18 29 06.4	+00 30 38	1.204 ± 0.121	6	0.32	2.07 ± 0.40	68	6.5^a	2.77 ± 0.14	0
12	18 29 09.1	+00 31 33	0.181 ± 0.024	–	–	–	75-A,B,C	1.7	1.16 ± 0.02	0
–	18 29 02.1	+00 31 21	< 0.0703	–	–	–	60	0.09	0.17 ± 0.02	I
–	18 29 09.1	+00 31 33	< 0.0904	–	–	–	63	0.12	0.22 ± 0.02	I
<i>DCE08-213:</i>										
13	18 29 52.2	+00 38 51.5	0.589 ± 0.067	–	–	–	–	–	–	–
14	18 29 53.3	+00 34 09.4	0.267 ± 0.082	–	–	–	–	–	–	–
–	18 29 53.04	+00 36 06.8	< 0.0950	–	–	–	157	0.57	–	–
–	18 29 52.52	+00 36 11.7	< 0.0950	–	–	–	154	1.0	0.54 ± 0.03	0
–	18 29 54.30	+00 36 01.3	< 0.0960	–	–	–	166	0.17	0.48 ± 0.05	I

^a From Harvey & Dunham (2009).

Table 6. AMI-LA observations of the Serpens filament. Columns are (1) AMI-LA source number; (2) right ascension of radio source peak; (3) declination of radio source peak; (4) integrated 1.8 cm flux density; (5) VLA association; (6) integrated 3.5 cm flux density; (7) spectral index from 3.5 to 1.8 cm; (8) associations from HH07; (9) bolometric luminosity from Enoch et al. (2009a), values in italics from EDJ09; (10) envelope mass from Enoch et al. (2009a) and (11) protostellar classification.

AMI	RA (J2000)	Dec. (J2000)	$S_{1.8\text{cm,int}}$ (mJy)	VLA	$S_{3.5\text{cm,int}}$ (mJy)	$\alpha_{3.5}^{1.8}$	[HH07]	L_{bol} (L_{\odot})	M_{env} (M_{\odot})	Cl.
<i>DCE08-202:</i>										
–	18 28 44.8	+00 51 26	< 0.095	–	–	–	24	0.04	0.24 ± 0.02	I
–	18 28 45.0	+00 52 04	< 0.098	–	–	–	26	1.09	0.29 ± 0.03	I
–	18 28 44.0	+00 53 38	< 0.138	–	–	–	23	0.19	0.26 ± 0.03	I

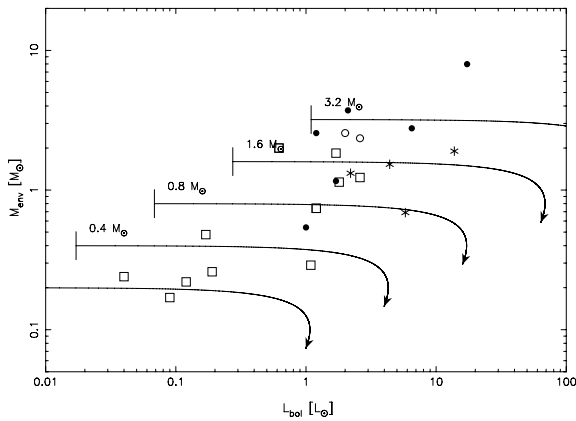


Figure 8. M_{env} versus L_{bol} diagram for Serpens. Class 0 protostars are shown as filled circles and Class I as open squares. Unclassified sources are shown as open circles. The four radio non-detections that are discrepant from the $L_{\text{rad}} - L_{\text{bol}}$ trend, but not the $L_{\text{rad}} - M_{\text{env}}$ trend, are shown as stars. Evolutionary tracks are shown for time-dependent accretion, $\dot{M}_{\text{acc}} \propto M_{\text{env}}(t)$, as described in the text, for initial mass reservoirs of $0.2 - 3.2 M_{\odot}$, as indicated in the figure.

the assumption that the accretion rate is proportional to the envelope mass, $\dot{M}_{\text{acc}} \propto M_{\text{env}}(t)$ (Bontemps et al. 1996; Motte & André 2001). In this scenario collapse is initiated within a finite reservoir of material with mass M_{env}^0 , with accretion proceeding

at a rate proportional to the mass of the residual envelope as a function of time, where $\dot{M}_{\text{acc}} = M_{\text{env}}/\tau$ and τ is a characteristic time-scale taken in this instance to be 10^5 yr. Following Bontemps et al. (1996), the envelope mass is assumed to be exponentially decreasing and consequently the accretion rate has the functional form $\dot{M}_{\text{acc}} = (M_{\text{env}}^0/\tau) \exp(-t/\tau)$, and the stellar mass: $M_* = (M_{\text{env}}^0)(1 - \exp(-t/\tau))$.

Under a simple standard stellar model (e.g. Adams, Lada & Shu 1987) we may assume that during the infall stage the bolometric luminosity is heavily dominated by accretion, $L_{\text{bol}} = L_{\text{acc}} = GM_*\dot{M}_{\text{acc}}/R_*$. We adopt a value of $R_* = 3 R_{\odot}$ (Stahler 1988), but note that for a more sophisticated model the stellar radius will in fact be a function of both the accretion rate and the stellar mass, as will the stellar luminosity, which will also contribute an additive term to the bolometric luminosity such that $L_{\text{bol}} = L_{\text{acc}} + L_*$. However in this simple scenario where we limit ourselves to the earlier stages of evolution we consider this term to be negligible. Models which are required to provide information past the characteristic time-scale will require these additional terms (see e.g. Motte & André 2001).

Compared to the evolutionary diagrams for the Taurus and ρ -Ophiucus molecular clouds (Motte & André 2001) the protostellar objects in Serpens seem to have evolved from initial mass reservoirs which are a factor of ≈ 2 larger. This apparent difference might be explained by the fact that the entire envelope will not be completely accreted on to the central star, but that a proportion of it will instead be dispersed by the act of outflows (Ladd, Fuller & Deane 1998). The four objects which are discrepant with the $L_{\text{rad}} - L_{\text{bol}}$ trend, but not

Table 7. Summary of detection statistics for Tables 4–6. The original sample includes only objects from DCE08, listed in Table 1. The extended sample includes all sources from E09, inclusive of the original sample.

Class	Present	Detected	per cent
<i>Original sample:</i>			
VeLLO	0	0	0/100
0	3	2	67
I	4	1	25
Group 1	1	1	100
Group 2	4	1	25
Group 3	3	0	0
<i>Extended sample:</i>			
0	6	4	67
I	17	3	18

the $L_{\text{rad}} - M_{\text{env}}$, have values of r which place them in a similar region of the evolutionary diagram. Under our simple scenario they appear to be objects which have evolved from a relatively large initial mass reservoir and are approaching their characteristic time-scale, where ≈ 50 per cent of their initial mass will have been accreted.

The detection statistics for all objects included in Tables 4–6 which lie within the FWHM of the AMI-LA primary beam and have an identified protostellar counterpart are summarized in Table 7. The small sky area, and consequently the small number of objects of each type present overall, means that no strong statistical conclusions can be drawn from this sample with the exception perhaps of the low Class I detection rate for the extended sample. However, we note that the detection rates for Class 0 and Class I objects in the extended sample are similar to those found for Perseus (Paper II) of 72 and 27 per cent, respectively.

6.2 Outflow momentum fluxes

Graves et al. (2010) calculate outflow momentum values for objects associated with four of the sub-mm cores in Serpens. Assuming a uniform dynamical age for each of 3×10^4 yr (Davis et al. 1999) we can estimate the momentum flux for these sources as $F_{\text{out}} = P/\tau$. The highly confused nature of this region precludes the use of more sophisticated momentum flux estimation methods. We use the maximum measured momentum from the blue and red outflow for each object, correcting the values in table 3 of Graves et al. (2010) by a factor of 2 to account for the full velocity range, as described in that paper; by a further factor of 10 to account for optical depth effects following Bontemps et al. (1996) and Visser et al. (2002); and by a small factor to account for the difference in distance assumed to Serpens. A comparison of these values with the overall correlation of radio luminosity to outflow momentum flux from Papers I and II is shown in Fig. 9. As already described in Papers I and II, following the model of Curiel et al. (1989) a significant proportion of objects exhibit radio luminosities which are inconsistent with free-free emission arising solely as a result of shock ionization in their molecular outflows. The four objects measured here (SMM 1, 4, 3/6 and 8) also appear to have radio emission inconsistent with this model, most markedly in the case of Serpens SMM 1. In Fig. 9 we show the limit set by the shock ionization model of Curiel et al. (1989) on radio luminosity for a canonical electron temperature of 10^4 K, typical of free-free emission, as a solid line. Also indicated is the same model for a temperature of 3000 K, the electron temperature indicated by *Herschel* observations of shocks

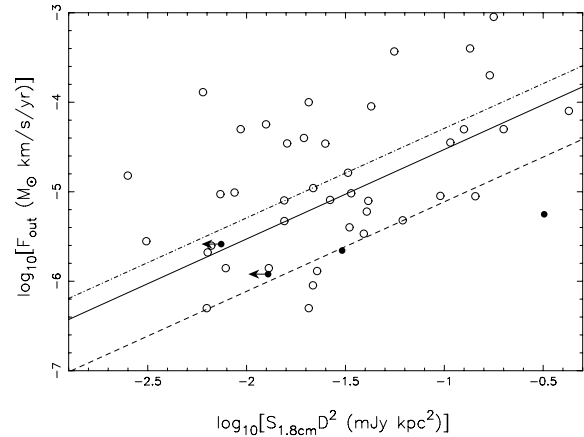


Figure 9. Correlation of radio luminosity with outflow momentum flux. Data from Papers I and II are shown as unfilled circles, and data from this work are shown as filled circles. The predicted relationship between momentum flux and radio luminosity (Curiel et al. 1989) for a 100 per cent efficiency and a temperature of 3000 K (dot-dashed line), 10^4 K (solid line) and 10^5 K is indicated.

(van Kempen et al. 2010), and a temperature of 10^5 K. It is evident that even for very high temperature electron gas the shock ionization model cannot account for the radio emission from a significant proportion of these objects.

However, the basis of a time-dependent accretion rate, such as that outlined in the simple exponential model of Section 6.1, is founded on the assumption that there is a relationship between outflow force and accretion rate. This assumption arises from the hypothesis that the energy required to provide the outflow is released through accretion on to the central protostar (Bontemps et al. 1996). If this is the case, and the observed correlation between radio luminosity, $S_\nu D^2$, and envelope mass shown in Fig. 7 is a true one, then there should be a corresponding relationship between outflow momentum and radio luminosity. Therefore it seems likely that the lack of a strong correlation in Fig. 9 is due to the inconsistent methods used to derive momentum flux across the sample.

6.3 Distance dependence

In the derived correlations and physical characteristics we have used a distance for the Serpens molecular cloud of 260 pc, consistent with the *Spitzer* catalogue. However, recently Dzib et al. (2010) have suggested that Serpens has a distance of 415 pc instead, using parallax measurements with the VLBA. This difference in assumed distance will affect the luminosity, mass and outflow momentum flux values quoted here. Radio luminosities will be raised by a factor of approximately 2.5, as will bolometric luminosities and envelope mass values, which are proportional to D^2 . Since each of these quantities is increased proportionately to the other no change will occur in the degree of correlation shown in Figs 6 and 7. Outflow momentum flux is linearly proportional to distance and, consequently, the outflow momentum flux values shown in Fig. 9 will be increased by a factor of 1.6; however the radio luminosity values will also be increased by a factor of 2.5. Since the increase in radio luminosity is proportionately greater than that of momentum flux this acts to worsen the agreement with the model (Curiel et al. 1989) further.

6.4 Radio spectral indices

The spectral index of partially optically thick free–free emission from fractionally ionized stellar winds and jets has been shown to vary depending on morphology (Reynolds 1986). Shallower spectral indices ($\alpha = 0.25$) are found for collimated outflows than in the spherical case ($0.6 < \alpha < 0.8$), and steeper spectra still are found for conical jets ($\alpha > 0.9$). In observations such as those presented here it is expected that unresolved contributions will be present in the radio spectrum from the core itself as well as a confined jet close to this core which then becomes unconfined at larger distances.

The mean value of spectral index for those sources which are detected at both 3.5 and 1.8 cm is $\bar{\alpha}_{3.5}^{1.8} = 0.53 \pm 1.14$, consistent with the canonical value for a partially optically thick spherical or collimated stellar wind of $\alpha = 0.6$. However the possibility of variable radio emission from these sources may limit the meaningfulness of this agreement. Certainly in the case of Serpens SMM 1 (DCE08-210), although we observe a non-thermal spectrum from 3.5 to 1.8 cm, previous observations at a similar epoch measured an optically thin free–free spectrum. Across the AMI-LA band we see a similar behaviour with $\alpha_{\text{AMI}} = -0.15 \pm 0.23$, consistent with optically thin free–free, although the short frequency coverage of the AMI-LA makes the uncertainty on this measurement large.

We do not fit spectral indices for AMI-7 and AMI-8 in the DCE08-215 field as the sources cannot be separated without the subtraction of AMI-8 from the map. This subtraction in itself will affect the spectrum of AMI-7. For the three radio detections in the DCE08-208 field we fit spectra across the AMI band; however the low signal-to-noise ratio for these objects, combined with the short frequency coverage of AMI, causes the uncertainties on these measurements to be large and consequently we suggest that they should only be used as an indication of emission mechanism rather than for quantitative comparisons.

HMH07-59 (AMI-10) has a falling spectral index across the AMI band, with $\alpha_{\text{AMI}} = -1.35 \pm 0.34$, differing from the optically thin index from 3.5 to 1.8 cm. There are a number of possible explanations for this difference, notably the 3.5 cm detection is associated solely with HMH07-59 whereas, although we also associate AMI-10 with this source, the 1.8 cm emission is extended and probably has contributions from multiple sources. The extension of the emission itself will cause a steepening of the spectrum across the AMI band, as even a small amount of flux loss over such a short frequency range can cause significant steepening. However, HMH07-59 is identified as a Class II/III object and therefore it is possible that the emission we are detecting is not thermal bremsstrahlung but gyrosynchrotron. Gyrosynchrotron emission is characterized by a sharply peaked spectrum, where the frequency of the peak is strongly dependent on the magnitude of the magnetic field, B . Following Dulk, Melrose & White (1979) the position of this peak can be determined by

$$\left[\nu_{\text{peak}} \text{GHz}^{-1} \right] \simeq 2.4 \left[\left(\frac{n_e}{10^{10} \text{cm}^{-3}} \right) \left(\frac{d}{10^8 \text{cm}} \right) \right]^{0.1} \times (\sin \theta)^{0.6} \left[\frac{T_e}{10^8 \text{K}} \right]^{0.7} \left[\frac{B}{10^2 \text{G}} \right]^{0.9},$$

where there is also a weak dependency on the electron density, n_e , and the path length, d , as well as the electron temperature of the plasma T_e and the inclination angle between the wave-normal direction and the magnetic field, θ . A flat spectral index from 3.5 to 1.8 cm and a falling spectral index across the AMI band imply that the peak of such a gyro-synchrotron spectrum would be found in the region of $\nu \simeq 12$ GHz. The electron density in such objects

is assumed to be in the range $10^3 \leq n_e \leq 10^9 \text{cm}^{-3}$, and the path length, roughly equivalent to the infall radius, to be in the range $0.01 \leq d \leq 0.1$ pc. Making the assumptions of $n_e = 10^5 \text{cm}^{-3}$ and $d = 0.1$ pc, as well as assuming an electron temperature of 10^8 K and an inclination angle of $\theta = \pi/4$, this would imply a magnetic field strength of $B = 270$ G, with approximately one order of magnitude uncertainty from the possible variation in parameters, as well as a small variation in the exponents dependent on the exact optical depth (Dulk et al. 1979).

Such a magnetic field is not unlikely in the vicinity of a Class II/III object. Typical magnetic field strengths of tens of Gauss have been suggested at a distance of a few astronomical units for objects with circumstellar discs (Shu et al. 2007), and much higher fields of a few kiloGauss have been measured from the surfaces of classical T Tauri stars (Johns-Krull, Valenti & Koresko 1999; Donati et al. 2005; Johns-Krull 2007). More recently, kiloGauss fields have also been measured from a Class I object (WL 17; Johns-Krull et al. 2009). A hypothesis of gyrosynchrotron emission from this object would be well tested by full radio polarization measurements, as gyrosynchrotron emission is known to be highly circularly polarized. As the AMI-LA measures I+Q it is unable to provide any information on the Stokes V component towards this object.

The remaining two sources, HMH07-68 (AMI-10) and HMH07-75 (AMI-11), appear to have optically thin spectra with $\alpha_{\text{AMI}} = -0.17 \pm 1.08$ and $\alpha_{\text{AMI}} = -0.05 \pm 1.15$, respectively. In the case of HMH07-68 this may indicate that the optical depth of the ionized medium has reached unity between 3.5 and 1.8 cm, typical of ultra-compact H II.

The conclusions drawn from comparisons of $\alpha_{3.5}^{1.8}$ and α_{AMI} are questionable as the 3.5 and 1.8 cm observations have been made at widely different epochs. The question of variability in the radio emission of protostars is still unanswered and any variation over the period intervening these measurements will make these comparisons invalid.

6.5 Evidence for radio variability in DCE08-210 (Serpens SMM 1)

Serpens SMM 1 was previously observed at 15 GHz by Snell & Bally (1986) who found a flux density of $S_{15} = 10 \pm 3$ mJy. Disregarding the large uncertainty in this measurement, it differs from the AMI-LA 16 GHz flux density by a factor of ≈ 2 . The 5 GHz flux density measured by Rodríguez et al. (1989) is $S_5 = 6.4 \pm 0.4$ mJy, a factor of 1.5 lower than that of Snell & Bally (1986). Additional evidence for variability may also be identified when comparing the flux densities from Snell & Bally (1986) with that at 8.5 GHz from Eiroa et al. (2005). Although the spectral index, $\alpha_5^{15} = 0.0 \pm 0.3$, measured by Snell & Bally (1986) from observations made at a similar epoch implies optically thin free–free emission, the 8.5 GHz flux density is discrepant from this model by $\approx 5\sigma$. Similarly the 25.4 GHz flux density of Ungerechts & Guesten (1984) of $S_{25.4} = 18 \pm 2$ mJy is significantly removed from the predictions of this model, in the opposite sense, by a degree which cannot be accounted for by an enhanced thermal dust contribution which is < 1 mJy.

Evidence for variable accretion in DCE08-210 has previously been suggested by Enoch et al. (2009b) who detect a high-mass disc around this source, which should be unstable and therefore exhibit luminosity variations. The radio data for this object suggest that there was previously a radio ‘flare’ event, the additional emission from which has now faded. We show this possibility in Fig. 10,

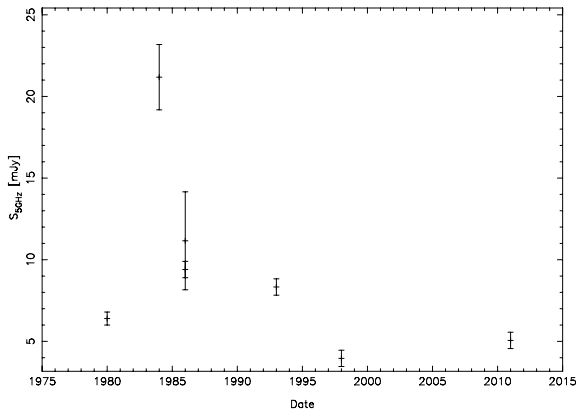


Figure 10. Radio flux density of Serpens SMM 1 adjusted to 5 GHz (see text for details) at different epochs.

where we have assumed that the spectrum of DCE08-210 is optically thin from 1.4 to 22 GHz in a single epoch and have corrected the measured flux densities to a common frequency of 5 GHz using a canonical spectral index for optically thin free-free emission of $\alpha = -0.1$.

7 CONCLUSIONS

The aim of this sample was to increase the data for Class I objects and improve the mass to radio luminosity correlation measured in Paper II. The low detection rate for Class I objects in Perseus is, however, echoed in Serpens and only 18 per cent of Class I objects present have been detected (three sources). Although this is only a small increase in the available data these objects do not demonstrate the same evolutionary divide hinted at in the correlation of radio flux density with envelope mass seen in Paper II. The objects detected in Serpens do, however, follow the general trends measured for radio luminosity with both bolometric luminosity and envelope mass, and the respective detection rates for Class 0 and Class I objects are consistent with those found for Perseus in Paper II. We also conclude that envelope mass provides a better radio detection indicator than bolometric luminosity with the deviations from the $L_{\text{rad}}-M_{\text{env}}$ correlation being notably fewer than that of the $L_{\text{rad}}-L_{\text{bol}}$ correlation.

ACKNOWLEDGEMENTS

We thank the staff of the Lord's Bridge Observatory for their invaluable assistance in the commissioning and operation of the Arcminute Microkelvin Imager. The AMI-LA is supported by Cambridge University and the STFC. CRG, TS, TF, MO and MLD acknowledge the support of PPARC/STFC studentships. AS would like to acknowledge support from Science Foundation Ireland under grant 07/RFP/PHYF790.

REFERENCES

Adams F. C., Lada C. J., Shu F. H., 1987, *ApJ*, 312, 788
 AMI Consortium: Hurley-Walker N. et al., 2009, *MNRAS*, 398, 249
 AMI Consortium: Scaife A. M. M. et al., 2008, *MNRAS*, 385, 809
 AMI Consortium: Scaife A. M. M. et al., 2011a, *MNRAS*, 415, 893 (Paper I)

AMI Consortium: Scaife A. M. M. et al., 2011b, *MNRAS*, 410, 2662 (Paper II)
 AMI Consortium: Zwart J. T. L. et al., 2008, *MNRAS*, 391, 1545
 André P., Motte F., Bacmann A., Belloche A., 1999, Nakamoto T., ed., *Proc. Star Formation*, p. 145
 Anglada G., 1995, *Rev. Mex. Astron. Astrofís.*, 1, 67
 Bontemps S., André P., Terebey S., Cabrit S., 1996, *A&A*, 311, 858
 Cambrésy L., 1999, in Bica M. D., Cutri R. M., Madore B. F., eds, *ASP Conf. Ser. Vol. 177, Astrophysics with Infrared Surveys: A Prelude to SIRTIF*. Astron. Soc. Pac., San Francisco, p. 256
 Casali M. M., Eiroa C., Duncan W. D., 1993, *A&A*, 275, 195
 Cohen M., Kuhl L. V., 1979, *ApJS*, 41, 743
 Condon J. J., Cotton W. D., Greisen E. W., Yin Q. F., Perley R. A., Taylor G. B., Broderick J. J., 1998, *AJ*, 115, 1693
 Curiel S., Rodríguez L. F., Bohigas J., Roth M., Canto J., Torrelles J. M., 1989, *Astrophys. Lett. Commun.*, 27, 299
 Curiel S., Rodríguez L. F., Moran J. M., Canto J., 1993, *ApJ*, 415, 191
 Davies M. L. et al., 2010, preprint (arXiv:1012.3659)
 Davis C. J., Matthews H. E., Ray T. P., Dent W. R. F., Richer J. S., 1999, *MNRAS*, 309, 141
 de Zotti G., Ricci R., Mesa D., Silva L., Mazzotta P., Toffolatti L., González-Nuevo J., 2005, *A&A*, 431, 893
 Di Francesco J., Johnstone D., Kirk H., MacKenzie T., Ledwosinska E., 2008, *ApJS*, 175, 277
 Djupvik A. A., André P., Bontemps S., Motte F., Olofsson G., Gälfalk M., Florén H.-G., 2006, *A&A*, 458, 789
 Donati J.-F., Paletou F., Bouvier J., Ferreira J., 2005, *Nat*, 438, 466
 Dulk G. A., Melrose D. B., White S. M., 1979, *ApJ*, 234, 1137
 Dullemond C. P., Dominik C., 2004, in Beaulieu J. P., Lecavelier des Etangs A., Terquem C., eds, *ASP Conf. Ser. Vol. 321, Extrasolar Planets: Today and Tomorrow*. Astron. Soc. Pac., San Francisco, p. 361
 Dunham M. M., Crapsi A., Evans N. J., II, Bourke T. L., Huard T. L., Myers P. C., Kauffmann J., 2008, *ApJS*, 179, 249 (DCE08)
 Dzib S., Loinard L., Mioduszewski A. J., Boden A. F., Rodríguez L. F., Torres R. M., 2010, *ApJ*, 718, 610
 Eiroa C., Casali M. M., 1992, *A&A*, 262, 468
 Eiroa C., Torrelles J. M., Curiel S., Djupvik A. A., 2005, *AJ*, 130, 643
 Enoch M. L., Glenn J., Evans N. J., II, Sargent A. I., Young K. E., Huard T. L., 2007, *ApJ*, 666, 982
 Enoch M. L., Corder S., Dunham M. M., Duchêne G., 2009a, *ApJ*, 707, 103
 Enoch M. L., Evans N. J., II, Sargent A. I., Glenn J., 2009b, *ApJ*, 692, 973
 Evans N. J. et al., 2009, *ApJS*, 181, 321 (EDJ09)
 Graves S. F. et al., 2010, *MNRAS*, 409, 1412
 Harvey P., Dunham M. M., 2009, *ApJ*, 695, 1495
 Harvey P. M. et al., 2006, *ApJ*, 644, 307
 Harvey P., Merín B., Huard T. L., Rebull L. M., Chapman N., Evans N. J., II, Myers P. C., 2007a, *ApJ*, 663, 1149 (HMH07)
 Harvey P. M. et al., 2007b, *ApJ*, 663, 1139
 Hatchell J., Fuller G. A., Richer J. S., Harries T. J., Ladd E. F., 2007, *A&A*, 468, 1009
 Hurt R. L., Barsony M., 1996, *ApJ*, 460, L45
 Johns-Krull C. M., 2007, *ApJ*, 664, 975
 Johns-Krull C. M., Valenti J. A., Koresko C., 1999, *ApJ*, 516, 900
 Johns-Krull C. M., Greene T. P., Doppmann G. W., Covey K. R., 2009, *ApJ*, 700, 1440
 Kaas A. A. et al., 2004, *A&A*, 421, 623
 Kenyon S. J., Hartmann L., 1995, *ApJS*, 101, 117
 Ladd E. F., Fuller G. A., Deane J. R., 1998, *ApJ*, 495, 871
 Motte F., André P., 2001, *A&A*, 365, 440
 Ossenkopf V., Henning T., 1994, *A&A*, 291, 943
 Panagia N., Felli M., 1975, *A&A*, 39, 1
 Patnaik A. R., Browne I. W. A., Wilkinson P. N., Wrobel J. M., 1992, *MNRAS*, 254, 655
 Reynolds S. P., 1986, *ApJ*, 304, 713
 Rodríguez L. F., Curiel S., Moran J. M., Mirabel I. F., Roth M., Garay G., 1989, *BAAS*, 21, 792

- Shu F. H., 1977, *ApJ*, 214, 488
Shu F. H., Galli D., Lizano S., Cai M. J., 2007, *IAU Symp.* 243, *Star Disk Interaction in Young Stars*. Kluwer, Dordrecht, p. 249
Smith K., Güdel M., Benz A. O., 1999, *A&A*, 349, 475
Snell R. L., Bally J., 1986, *ApJ*, 303, 683
Stahler S. W., 1988, *BAAS*, 20, 1034
Straižys V., Černis K., Bartašiūtė S., 1996, *Baltic Astron.*, 5, 125
Ungerechts H., Guesten R., 1984, *A&A*, 131, 177
van Kempen T. A. et al., 2010, *A&A*, 518, L121
Visser A. E., Richer J. S., Chandler C. J., 2002, *AJ*, 124, 2756
Winston E. et al., 2007, *ApJ*, 669, 493
Young C. H., Evans N. J., II, 2005, *ApJ*, 627, 293
Young C. H. et al., 2004, *ApJS*, 154, 396

This paper has been typeset from a \TeX/L\AA\TeX file prepared by the author.