

A DISTANT RADIO MINI-HALO IN THE PHOENIX GALAXY CLUSTER

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Draft version September 25, 2018

ABSTRACT

We report the discovery of extended radio emission in the Phoenix cluster (SPT-CL J2344–4243, $z=0.596$) with the GMRT at 610 MHz. The diffuse emission extends over a region of at least 400–500 kpc and surrounds the central radio source of the Brightest Cluster Galaxy, but does not appear to be directly associated with it. We classify the diffuse emission as a radio mini-halo, making it the currently most distant mini-halo known. Radio mini-halos have been explained by synchrotron emitting particles re-accelerated via turbulence, possibly induced by gas sloshing generated from a minor merger event. Chandra observations show a non-concentric X-ray surface brightness distribution, which is consistent with this sloshing interpretation. The mini-halo has a flux density of 17 ± 5 mJy, resulting in a 1.4 GHz radio power of $(10.4 \pm 3.5) \times 10^{24}$ W Hz⁻¹. The combined cluster emission, which includes the central compact radio source, is also detected in a shallow GMRT 156 MHz observation and together with the 610 MHz data we compute a spectral index of -0.84 ± 0.12 for the overall cluster radio emission. Given that mini-halos typically have steeper radio spectra than cluster radio galaxies, this spectral index should be taken as an upper limit for the mini-halo.

Subject headings: Galaxies: clusters: individual (SPT-CL J2344–4243) — Galaxies: clusters: intracluster medium — large-scale structure of universe — Radiation mechanisms: non-thermal — X-rays: galaxies: clusters

1. INTRODUCTION

Radio mini-halos are diffuse extended sources that are found in some relaxed, cool-core clusters (see the review by Feretti et al. 2012). They have steep radio spectra ($\alpha < -1$, with $S_\nu \propto \nu^\alpha$) and sizes of $\sim 100 - 500$ kpc, somewhat smaller than Mpc-scale giant radio halos that are found in merging galaxy clusters. The emission from mini-halos typically surrounds the central radio source associated with the Brightest Cluster Galaxy (BCG). Mini-halos are rare with less than 20 confirmed cases (e.g., Feretti et al. 2012; Giacintucci et al. 2014b). The prototypical example is the mini-halo found in the Perseus cluster (Sijbring 1993; Gitti et al. 2002).

Considering the lifetimes of the radio emitting electrons and the extent of the mini-halos, a form of in-situ cosmic ray (CR) production or re-acceleration is required to explain their presence. Several different possibilities have been put forward. Hadronic (also called secondary) models invoke collisions between CR protons and thermal protons to produce a population of secondary relativistic electrons (Pfrommer & Enßlin 2004; Fujita et al. 2007; Keshet & Loeb 2010; Fujita & Ohira 2013). Alternatively, the synchrotron emitting electrons could arise from mildly relativistic electrons (for example cooled electrons from the central radio galaxy) that

have been re-accelerated by turbulence (Gitti et al. 2002, 2004).

Recently, an interesting correspondence between cold fronts and mini-halos has been found for some clusters. Based on this, the idea has been put forward that sloshing motions, induced by minor mergers, generate turbulence in the cores of clusters, which is then capable of re-accelerating electrons. In the presence of a magnetic field, these electrons emit synchrotron emission and could form a mini-halo (Mazzotta & Giacintucci 2008; ZuHone et al. 2013). Brunetti & Jones (2014) also mention the possibility of a more direct connection and possible transition between mini-halos and giant radio halos in clusters.

In this letter we report on the results of Giant Metrewave Radio Telescope (GMRT) radio observations of the Phoenix cluster (SPT-CL J2344–4243), which were taken as part of a larger radio survey of distant massive galaxy clusters (i.e., van Weeren et al. 2014). The Phoenix cluster was discovered from its Sunyaev-Zel’dovich effect signal with the South Pole Telescope (Williamson et al. 2011) and is located at $z = 0.596 \pm 0.002$ (McDonald et al. 2012). The BCG appears to be undergoing a strong starburst with a star formation rate of $750 \pm 160 M_\odot \text{ yr}^{-1}$ (McDonald et al. 2012). This cluster is the most extreme cool-core cluster known ($\dot{M} \approx 2700 \pm 700 M_\odot \text{ yr}^{-1}$; McDonald et al. 2013) and is also very massive ($M_{500} \approx 1.3 \times 10^{15} M_\odot$). Inter-

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estingly, the combination of a high cooling rate and the relatively weak central radio source (associated with the BCG) suggests that feedback has been unable to halt cooling in this system, leading to the high star formation rate.

In this letter we adopt a Λ CDM cosmology with $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.27$, and $\Omega_\Lambda = 0.73$. With the adopted cosmology, $1''$ corresponds to a physical scale of 6.652 kpc at $z = 0.596$. The resulting luminosity distance (D_L) is 3495 Mpc. All our images are in the J2000 coordinate system.

2. OBSERVATIONS & DATA REDUCTION

SPT-CL J2344–4243 was observed with the GMRT on June 14 and 15, 2013, for a total of about 10 hrs. Due to the low-elevation of the cluster at the GMRT site, the observing time was divided into two separate sessions. The continuum observations were carried out at a frequency of 610 MHz with a usable bandwidth of 29 MHz. In addition, we used GMRT 156 MHz archival data taken on August 5, 2011, with a total on source time of 13 min. Both datasets were reduced via a semi-automatic data reduction scheme (Intema et al. 2009) employing the Astronomical Image Processing System² (AIPS), ParselTongue (Kettenis et al. 2006) and Orbit (Cotton 2008). The data reduction started with automatic radio frequency interference (RFI) removal, band-pass and gain calibration. The flux calibration was performed according to the scale described in Scaife & Heald (2012). The calibration was further refined via several cycles of self-calibration. Low-level broadband RFI was subtracted using an improved version of the method described by Athreya (2009). After self-calibration, direction-dependent calibration solutions were obtained for sources with a signal to noise ratio $> 150 - 200$. The imaging was done via facets to apply the direction-dependent calibration solutions and to correct for the non-coplanar nature of the array.

Archival data from a *Chandra X-ray Observatory* ACIS-I observation were also studied. The 11.9 ks of data (ObsId 13401) were calibrated following the processing described in Vikhlinin et al. (2005), applying the most recent calibration files³. The calibration included the application of gain maps to calibrate photon energies, a correction for position-dependent charge transfer inefficiency, and filtering of counts with a recomputed ASCA grade 1, 5, or 7 and those from bad pixels. In addition, periods of elevated background were filtered by examining the light curves. Periods with count rates a factor of 1.2 above the mean count rate in the 6–12 keV band were also removed. The discarded exposure time was negligible (132 s). The final image was made in the 0.5–2.0 keV band to increase the contrast between the thermal Bremsstrahlung and AGN emission from the central BCG, and to maximize the ratio between source and background counts. We used a pixel binning factor of 2.

3. RESULTS

Our 610 MHz image of the Phoenix cluster, with a resolution of $13.7'' \times 5.8''$, is shown in Fig. 1 (left

TABLE 1
CLUSTER AND RADIO MINI-HALO PROPERTIES

z	0.596 ± 0.002
R_{500} (Mpc)	1.3
$L_{X,500}$ ($10^{44} \text{ erg s}^{-1}$, 2.0–10.0 keV)	82_{-2}^{+1}
T_{500} (keV)	$13.0_{-3.4}^{+2.5}$
M_{500} ($10^{15} M_\odot$)	$1.26_{-1.5}^{+2.0}$
$S_{\text{mini-halo},610\text{MHz}}$ (mJy)	17 ± 5
$S_{4\sigma_{\text{rms}} \text{ enclosed},610\text{MHz}}$ (mJy)	114 ± 11^a
α_{156}^{610} (compact and diffuse emission)	-0.84 ± 0.12
$P_{\text{mini-halo},1.4\text{GHz}}$ ($10^{24} \text{ W Hz}^{-1}$)	10.4 ± 3.5^b
LLS _{mini-halo} (Mpc)	0.4–0.5
610 MHz compact source fluxes (mJy)	$89.7^c \pm 9.0$ (A)
	4.29 ± 0.44 (B)
	3.46 ± 0.36 (C)
	3.95 ± 0.41 (D)
	1.11 ± 0.15 (E)

Cluster properties taken from McDonald et al. (2012)

^a integrated flux density reported by the AIPS task TVSTAT for all emission enclosed within the $4\sigma_{\text{rms}}$ contours of Fig. 1 (left panel), this includes sources A, B, C and the diffuse emission

^b assuming a spectral index of $\alpha = -1.1$ for the mini-halo emission

^c source A has a peak flux of 70.5 ± 7.1 mJy

panel). It reveals a central compact source associated with the BCG. This central source is also detected in the 843 MHz Sydney University Molonglo Sky Survey (SUMSS, Bock et al. 1999; Mauch et al. 2003) with a flux density of $S_{843} = 79.2 \pm 2.8$ mJy. The 610 MHz image shows four additional sources within about $1'$ of the BCG. The image also reveals extended diffuse emission surrounding the central compact radio source. Subtracting the emission from compact sources and tapering down to $25''$ resolution, we find a total extent of ~ 0.8 to $\sim 1.2'$, with $1.2'$ corresponding to a largest linear size (LLS) of 480 kpc. This low-resolution image is shown in Fig. 2 (right), more details are given later in the Section.

The global cluster properties are summarized in Table 1. A higher resolution image ($9.4'' \times 3.6''$) to emphasize compact sources was made by excluding baselines below $7 \text{ k}\lambda$ (corresponding to a spatial scale of about $0.6'$) and is shown in the right panel of Fig. 1. An overlay on an archival Spitzer 4.5 micron IRAC image is shown in Fig. 2 with compact radio sources labeled. It reveals counterparts to most of the radio sources visible in the high-resolution image. The GMRT 156 MHz image is overlaid in Fig. 1 (right panel). This shallow image only displays the central compact radio source, but there is a hint of a southward extension towards compact source C.

Because of our limited spatial resolution and the fact that the cluster is located at $z = 0.596$, it is somewhat difficult to accurately separate the diffuse emission from the central compact radio source. In addition, the contribution from other compact sources needs to be subtracted. The integrated flux densities of the compact sources were measured with the AIPS task JMFIT, fitting Gaussian components to the sources. The resulting integrated flux densities are given in Table 1. The quoted uncertainties are based on the map noise and an absolute flux calibration uncertainty of 10% (e.g., Chandra et al. 2004).

The presence of diffuse emission on scales of $\sim 400 - 500$ kpc at 610 MHz, provides the first evidence that the central region of the Phoenix cluster contains an ex-

² <http://www.aips.nrao.edu>

³ We used CIAO 4.6 and CALDB 4.5.9

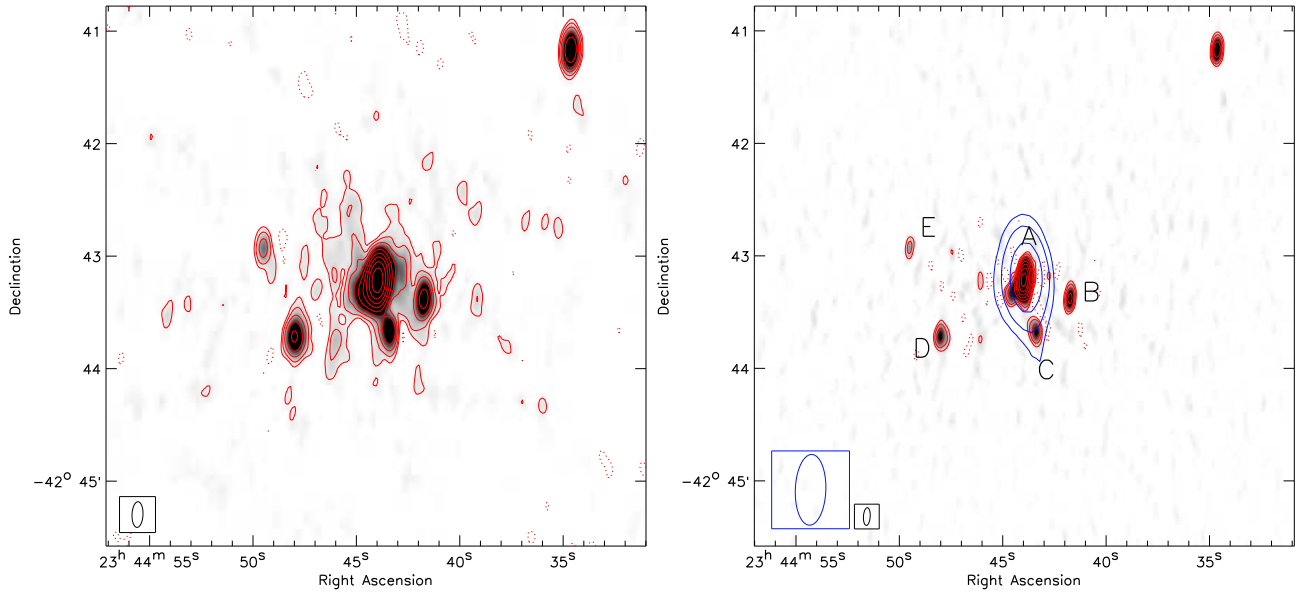


FIG. 1.— Left: GMRT 610 MHz image of the Phoenix cluster. The image was made using Briggs weighing (Briggs 1995) with the robust parameter set to 0.5. Contour levels are drawn at $[1, 2, 4, 8, \dots] \times 4\sigma_{\text{rms}}$, with $\sigma_{\text{rms}} = 43 \mu\text{Jy beam}^{-1}$. Negative $-3\sigma_{\text{rms}}$ contours are shown with dotted lines. The beam size is $13.7'' \times 5.8''$ and indicated in the bottom left corner. Right: High-resolution ($9.4'' \times 3.6''$) GMRT 610 MHz image is shown with red contours (and in grayscale). The contours are drawn at the same levels as in the left panel (but with $\sigma_{\text{rms}} = 95 \mu\text{Jy beam}^{-1}$). Compact sources are labelled. For the imaging, data below $7 \text{ k}\lambda$ were excluded and the robust parameter was set to 0.0. The GMRT 156 MHz image is shown with blue contours. The image has a central noise level of 8 mJy beam^{-1} and the contour levels are drawn at $[1, 2, 4, 8, \dots] \times 40 \text{ mJy beam}^{-1}$. The beam size is $38'' \times 16''$.

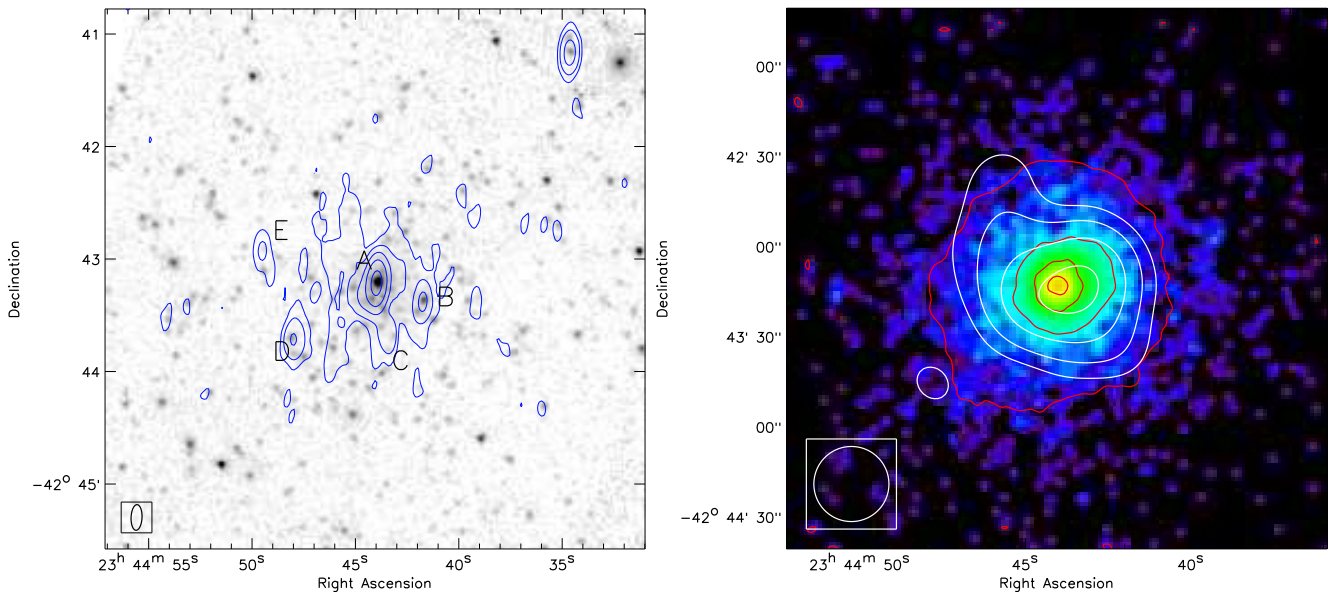


FIG. 2.— Left: Spitzer 4.5 micron IRAC image overlaid with GMRT 610 MHz contours drawn at $[1, 4, 16, \dots] \times 4\sigma_{\text{rms}}$. Contours are from the image shown in Fig. 1 (left panel). Right: Chandra 0.5–2.0 keV ACIS-I image of the Phoenix cluster smoothed with a $3''$ FWHM Gaussian (pixel size is $0.984''$). Red contours are X-ray isophotes drawn at levels of $[1, 10, 50, 200] \times 10^{-5} \text{ cts s}^{-1} \text{ pixel}^{-1}$. GMRT 610 MHz low-resolution radio contours are overlaid in white. Compact sources (Fig. 1; right panel) were subtracted from the uv-data prior to making this image. Contour levels are drawn at $\sqrt{[1, 2, 4, 8, \dots]} \times 0.6 \text{ mJy beam}^{-1}$. The circular beam size is $25''$ and indicated in the bottom left corner.

tended non-thermal component. In principle two possible explanations can be given for the diffuse component. The diffuse emission either originates from the radio lobes of the central AGN (source A) or it indicates the existence of a mini-halo. Since the diffuse emission surrounds the central AGN and its morphology does not resemble that of radio lobes, we classify this extended emission as a radio mini-halo. The physical extent of this mini-halo is also very similar to that of other known mini-halos in cool-core clusters (e.g., Cassano et al. 2008). A small extension towards the SE of source A is seen in the high-resolution image (Fig. 1; right panel). We attribute this emission to source A, and not to the mini-halo. Higher resolution radio observations are required to determine the precise nature of this extension.

We estimate the integrated mini-halo flux density using two different methods. For the first method, we simply integrate the flux enclosed within the $4\sigma_{\text{rms}}$ contours from Fig. 1 (left panel) and subtract the flux densities from the compact sources (A, B, and C) as are given in Table 1. In this way we obtain a mini-halo flux density of 16.6 ± 1.7 mJy. For the second method, we did attempt to remove the flux from the compact sources directly from the uv-data by subtracting the clean components from the high-resolution image shown in Fig 1 (right panel). We then re-imaged the data tapering to a resolution of $25''$ resolution. This image is shown overlaid on the Chandra 0.5–2.0 keV X-ray image in Fig. 2. The diffuse radio emission peaks approximately at the same position as the X-ray emission. From this low-resolution image we obtain a flux of 9 ± 1 mJy. This value likely underestimates the mini-halo flux density because some emission from the mini-halo has been subtracted since the mini-halo contributes flux on scales larger than 7 k λ . If we double the inner uv-range cut for making the high-resolution clean component model, we obtain a flux of 19 ± 2 mJy for the radio mini-halo.

For the remainder of this Letter we adopt a flux density of 17 ± 5 mJy for the mini-halo. This was obtained by integrating within the $4\sigma_{\text{rms}}$ contours and subtracting the compact sources flux densities, but with the uncertainty of 5 mJy reflecting the difficulties involved in separating the diffuse emission from the compact sources (i.e., approximately covering the range between the highest and lowest flux estimates reported in the previous paragraph). With this integrated flux density, we calculate a k -corrected 1.4 GHz radio power of $P_{1.4\text{GHz}} = (10.4 \pm 3.5) \times 10^{24}$ W Hz $^{-1}$ using $P_{1.4\text{GHz}} = 4\pi S_{1.4\text{GHz}} D_L^2 (1+z)^{-(\alpha+1)}$. For the mini-halo spectral index we adopted a value of $\alpha = -1.1$ (e.g., Giacintucci et al. 2014a,b; Murgia et al. 2010)

At 156 MHz we measure a flux density 330 ± 45 mJy for the combined cluster emission (diffuse emission + compact sources), in a region twice the 156 MHz beam size centered on the peak emission. Taking precisely the same region at 610 MHz, we compute an overall spectral index (central AGN and mini-halo combined) of -0.84 ± 0.12 . Typically mini-halos have steeper radio spectra than the central AGN and therefore the mini-halo emission could be considerably steeper than -0.84 . In particular because the central compact core dominates the integrated flux density. However, resolved multi-frequency observations are needed to obtain separate spectral indices for

the compact and diffuse components. In addition, deeper observations are needed to determine the full angular extent of the mini-halo.

A closer inspection of the Chandra image reveals non-concentric X-ray isophotes (Fig 2). The X-ray emission is more extended in the NW direction at $\sim 20''$ from the cluster center, while at a radius of $\sim 1'$ the cluster emission extends more in the SE direction.

4. DISCUSSION AND CONCLUSIONS

We report the discovery of a 400 – 500 kpc (total extent) radio mini-halo in the Phoenix cluster using GMRT 610 MHz observations. At a redshift of 0.596 ± 0.002 , the mini-halo in the Phoenix cluster is currently the most distant one known. Combining the 610 MHz data with shallow 156 MHz GMRT observations, we estimate an upper limit of -0.84 ± 0.12 for the mini-halo spectral index. By assuming a spectral index of -1.1 we estimated that the mini-halo has a 1.4 GHz radio power of $(10.4 \pm 3.5) \times 10^{24}$ W Hz $^{-1}$, higher than the Perseus cluster mini-halo (Sijbring 1993) but less than the RX J1347.5–1145 mini-halo (Gitti et al. 2007). This radio power/luminosity falls in the range that is expected for mini-halos (e.g., Giacintucci et al. 2014b).

Gitti et al. (2004, 2007) suggested a connection between the origin of the mini-halo synchrotron emission and cooling flows in clusters. This should result in a correlation between $P_{1.4\text{GHz}}$ and the cooling flow power: $P_{\text{CF}} = \dot{M}kT/\mu m_p$. Such a correlation is observed (e.g., Doria et al. 2012; Gitti et al. 2012). For the Phoenix cluster we have $P_{\text{CF}} = 3.5 \times 10^{45}$ erg s $^{-1}$. Given this extremely high cooling flow power, we would expect the Phoenix cluster mini-halo to have the highest known $P_{1.4\text{GHz}}$ of any cluster, which is however not the case. However, it is important to note that the observed scatter is quite large for the correlation and also our derived cooling power is uncertain due to the uncertain mass inflow rate ($\dot{M} \approx 2700 \pm 700 M_{\odot}$ yr $^{-1}$). More accurate measurements of \dot{M} and also $P_{1.4\text{GHz}}$ will be needed to determine whether the Phoenix cluster mini-halo deviates from the $P_{1.4\text{GHz}}-P_{\text{CF}}$ correlation.

The presence of a mini-halo in the Phoenix cluster has some interesting implications, related to the association of mini-halos, cold fronts and gas sloshing. Simulations show that cold fronts may be caused by gas sloshing, induced by a minor merger event (e.g., Ascasibar & Markevitch 2006). Minor mergers with a non-zero impact parameter result in cold fronts that form a non-concentric spiral-like pattern in the X-ray gas. These spiral-type patterns have been observed in cool-core clusters and many of these clusters also host mini-halos (e.g., Johnson et al. 2010; Laganá et al. 2010). In a few cases, the radio emission displays a spatial correlation with the spiral-like patterns in the X-ray gas (Mazzotta & Giacintucci 2008). This indicates that the synchrotron emitting electrons are trapped in the same ICM flows that resulted in the spiral-like pattern. Since minor merger events also inject turbulence into cluster cores, it has been proposed that this turbulence could re-accelerate a population of fossil electrons (for example from the central radio galaxy) giving rise to diffuse radio emission in the form of a radio mini-halo (see the numerical simulations by ZuHone et al. 2013).

In the above framework, the presence of a mini-halo in the Phoenix cluster implies that the cluster core has enough turbulent energy, possibly caused by sloshing gas motions, to re-accelerate a population of relativistic particles. This interpretation is also consistent with the non-concentric X-ray isophotes we found. Since the gas-sloshing motions could be induced by a merger event, this raises the question whether the very high star formation rate measured for the Phoenix Cluster core could (partly) be the result of a merger event with a gas-rich system. This possibility is however disfavored by McDonald et al. (2013). Upcoming Chandra observations might have the sensitivity to detect cold fronts and spiral-like patterns to determine if the mini-halo in the Phoenix could be related to sloshing motions and turbulence.

Acknowledgments: We would like to thank the anonymous referee for useful comments and the staff of the GMRT who have made these observations possible. The GMRT is run by the National Centre for Radio Astrophysics of the Tata Institute of Fundamental Research. This work is based in part on observations made with

the Spitzer Space Telescope, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA.

R.J.W. is supported by NASA through the Einstein Postdoctoral grant number PF2-130104 awarded by the Chandra X-ray Center, which is operated by the Smithsonian Astrophysical Observatory for NASA under contract NAS8-03060. H.T.I acknowledges support from the National Radio Astronomy Observatory, which is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc. M.B and M.H. acknowledge support by the research group FOR 1254 funded by the Deutsche Forschungsgemeinschaft: “Magnetisation of interstellar and intergalactic media: the prospects of low-frequency radio observations”. S.E.N. is supported by the DFG NU 332/2-1. W.R.F., C.J., and F.A-S. acknowledge support from the Smithsonian Institution. We thank Annalisa Bonafede and Julie Hlavacek-Larrondo for commenting on an earlier version of this letter.

Facilities: GMRT, CXO

REFERENCES

- Ascasibar, Y., & Markevitch, M. 2006, *ApJ*, 650, 102
 Athreya, R. 2009, *ApJ*, 696, 885
 Bock, D. C.-J., Large, M. I., & Sadler, E. M. 1999, *AJ*, 117, 1578
 Briggs, D. S. 1995, PhD thesis, New Mexico Institute of Mining Technology, Socorro, New Mexico, USA
 Brunetti, G., & Jones, T. W. 2014, *ArXiv e-prints*
 Cassano, R., Gitti, M., & Brunetti, G. 2008, *A&A*, 486, L31
 Chandra, P., Ray, A., & Bhatnagar, S. 2004, *ApJ*, 612, 974
 Cotton, W. D. 2008, *PASP*, 120, 439
 Doria, A., Gitti, M., Etori, S., et al. 2012, *ApJ*, 753, 47
 Feretti, L., Giovannini, G., Govoni, F., & Murgia, M. 2012, *A&A Rev.*, 20, 54
 Fujita, Y., Kohri, K., Yamazaki, R., & Kino, M. 2007, *ApJ*, 663, L61
 Fujita, Y., & Ohira, Y. 2013, *MNRAS*, 428, 599
 Giacintucci, S., Markevitch, M., Brunetti, G., et al. 2014a, *ArXiv e-prints*
 Giacintucci, S., Markevitch, M., Venturi, T., et al. 2014b, *ApJ*, 781, 9
 Gitti, M., Brighenti, F., & McNamara, B. R. 2012, *Advances in Astronomy*, 2012
 Gitti, M., Brunetti, G., Feretti, L., & Setti, G. 2004, *A&A*, 417, 1
 Gitti, M., Brunetti, G., & Setti, G. 2002, *A&A*, 386, 456
 Gitti, M., Ferrari, C., Domainko, W., Feretti, L., & Schindler, S. 2007, *A&A*, 470, L25
 Intema, H. T., van der Tol, S., Cotton, W. D., et al. 2009, *A&A*, 501, 1185
 Johnson, R. E., Markevitch, M., Wegner, G. A., Jones, C., & Forman, W. R. 2010, *ApJ*, 710, 1776
 Keshet, U., & Loeb, A. 2010, *ApJ*, 722, 737
 Kettenis, M., van Langevelde, H. J., Reynolds, C., & Cotton, B. 2006, in *Astronomical Society of the Pacific Conference Series*, Vol. 351, *Astronomical Data Analysis Software and Systems XV*, ed. C. Gabriel, C. Arviset, D. Ponz, & S. Enrique, 497
 Laganá, T. F., Andrade-Santos, F., & Lima Neto, G. B. 2010, *A&A*, 511, A15
 Mauch, T., Murphy, T., Buttery, H. J., et al. 2003, *MNRAS*, 342, 1117
 Mazzotta, P., & Giacintucci, S. 2008, *ApJ*, 675, L9
 McDonald, M., Benson, B., Veilleux, S., Bautz, M. W., & Reichardt, C. L. 2013, *ApJ*, 765, L37
 McDonald, M., Bayliss, M., Benson, B. A., et al. 2012, *Nature*, 488, 349
 Murgia, M., Govoni, F., Feretti, L., & Giovannini, G. 2010, *A&A*, 509, A86
 Pfrommer, C., & Enßlin, T. A. 2004, *A&A*, 413, 17
 Scaife, A. M. M., & Heald, G. H. 2012, *MNRAS*, 423, L30
 Sijbring, L. G. 1993, PhD thesis, University of Groningen
 van Weeren, R. J., Intema, H. T., Lal, D. V., et al. 2014, *ApJ*, 781, L32
 Vikhlinin, A., Markevitch, M., Murray, S. S., et al. 2005, *ApJ*, 628, 655
 Williamson, R., Benson, B. A., High, F. W., et al. 2011, *ApJ*, 738, 139
 ZuHone, J. A., Markevitch, M., Brunetti, G., & Giacintucci, S. 2013, *ApJ*, 762, 78