

Searching for pulsars associated with the *Fermi* GeV excess

D. Bhakta,^{1,2} J. S. Deneva,^{3★} D. A. Frail,^{2★} F. de Gasperin,⁴ H. T. Intema,⁴
P. Jagannathan^{2,5} and K. P. Mooley^{6†}

¹Department of Physics, Texas Tech University, Box 41051, Lubbock, TX 79409-1051, USA

²National Radio Astronomy Observatory, 1003 Lopezville Road, Socorro, NM 87801, USA

³Resident at Naval Research Laboratory, George Mason University, Washington, DC 20375, USA

⁴Leiden Observatory, Leiden University, Niels Bohrweg 2, NL-2333 CA Leiden, the Netherlands

⁵Department of Astronomy, University of Cape Town, Private Bag X3, Rondebosch 7701, Republic of South Africa

⁶Astrophysics, Department of Physics, University of Oxford, Keble Road, Oxford OX1 3RH, UK

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ABSTRACT

The *Fermi* Large Area Telescope has detected an extended region of GeV emission towards the Galactic Centre that is currently thought to be powered by dark matter annihilation or a population of young and/or millisecond pulsars. In a test of the pulsar hypothesis, we have carried out an initial search of a 20 deg² area centred on the peak of the galactic centre GeV excess. Candidate pulsars were identified as a compact, steep spectrum continuum radio source on interferometric images and followed with targeted single-dish pulsation searches. We report the discovery of the recycled pulsar PSR 1751–2737 with a spin period of 2.23 ms. PSR 1751–2737 appears to be an isolated recycled pulsar located within the disc of our Galaxy, and it is not part of the putative bulge population of pulsars that are thought to be responsible for the excess GeV emission. However, our initial success in this small pilot survey suggests that this hybrid method (i.e. wide-field interferometric imaging followed up with single-dish pulsation searches) may be an efficient alternative strategy for testing whether a putative bulge population of pulsars is responsible for the GeV excess.

Key words: surveys – stars: neutron – pulsars: general – pulsars: individual: PSR J1751–2737 – gamma-rays: general – radio continuum: general.

1 INTRODUCTION

Observations by the *Fermi* Large Area Telescope (LAT) have revealed an excess of GeV emission from the centre of our Galaxy (see review by van Eldik 2015). This diffuse emission can be seen over a 10° × 10° region towards the galactic centre but is strongest within a 2.5 radius of Sgr A* (Ajello et al. 2016). There are currently two alternative explanations that have been offered to explain this diffuse excess: (a) it is the long-sought annihilation signature of dark matter particles (Vitale et al. 2009; Hooper & Goodenough 2011; Weniger 2012; Daylan et al. 2016) or (b) the integrated high energy emission from a population of several thousand young and/or millisecond pulsars (Abazajian 2011; Brandt & Kocsis 2015). Recent analysis of the spatial and spectral properties of the gamma-ray excess strongly favours the pulsar hypothesis (Abazajian 2011; Mirabal 2013; Yuan & Zhang 2014; Calore, Cholis & Weniger 2015; Ajello et al. 2016; Bartels, Krishnamurthy & Weniger 2016; Lee et al. 2016).

The essential test of the pulsar hypothesis would be to detect individual bulge pulsars and show that their properties are consistent with the GeV excess signal. It has long been argued from theoretical grounds and multiwavelength observations that a substantial population of pulsars exists both in close orbit around the black hole Sgr A* and on larger scales of hundreds of parsecs around the galactic centre (Pfahl & Loeb 2004; Muno et al. 2005; Wang, Lu & Gotthelf 2006; Wharton et al. 2012). However, despite extensive searches for radio pulsations, only a handful of normal (i.e. non-recycled) pulsars and one magnetar have been discovered within a few hundred parsecs of the galactic centre, but no millisecond pulsars (Johnston et al. 2006; Deneva, Cordes & Lazio 2009; Eatough et al. 2013; Mori et al. 2013; Shannon & Johnston 2013).

The observational challenges in finding the putative bulge pulsars responsible for the gamma-ray excess have recently been summarized by Calore et al. (2016) and O’Leary et al. (2016). The large distance and the high brightness temperature of the diffuse synchrotron emission towards the bulge means that any pulsed signals will be weak and hard to detect without deep integrations. More problematic, however, is the large amount of ionized gas at the galactic centre, which results in heavy scattering and dispersion broadening (Frail et al. 1994; Lazio & Cordes 1998b; Macquart &

* E-mail: julia.deneva@gmail.com (JSD); dfrail@nrao.edu (DAF)

† Hintze Research Fellow.

Kanekar 2015). The effects are strongest within the central 150 pc radius of Sgr A*, but dispersive smearing and scattering remain large over most of the area of the GeV excess (see figs 5 and 6 of Calore et al. 2016). Temporal smearing of pulsed signals make them hard to detect unless observing is carried out at high frequencies ($\nu > 5$ GHz) where temporal scattering is sharply reduced ($\tau_{\text{scat}} \propto \nu^{-4}$). This approach has been used for targeted radio pulsar searches in the immediate vicinity of SgrA* (e.g. Macquart et al. 2010; Eatough et al. 2013; Siemion et al. 2013) but it is not feasible today to carry out a blind pulsation survey over the entire region of the GeV excess.

An alternative, potentially more efficient, approach is to begin by initially identifying candidates in the image plane (Cordes & Lazio 1997; Lazio & Cordes 2008) and then later follow it up with deep radio pulsations searches at higher frequencies. Pulsar candidates can be recognized in the image plane as compact, steep spectrum (polarized) radio sources. This image-based approach has been used in the past to find pulsars with some success (e.g. Navarro et al. 1995; Broderick et al. 2016; Frail et al. 2016a). In this paper, we carry out a pilot study using the recently published GMRT 150 MHz all-sky survey (TGSS ADR; Intema et al. 2017) and existing images at higher frequencies to search for compact, steep spectrum radio sources within the central 2:5 radius of the gamma-ray excess. In Section 2, we describe our candidate selection method; Section 3 summarizes our observations of two candidates, data processing and results; and Section 4 presents conclusions.

2 METHOD

We generated our initial candidate list from a recently published continuum all-sky survey carried out on the Giant Metrewave Radio Telescope (GMRT) at a frequency of 150 MHz (TGSS ADR; Intema et al. 2017). Pulsars stand out at 150 MHz. With their steep, power-law spectra, the true distribution spectral indices have a mean $\alpha = -1.4 \pm 1$ (Bates, Lorimer & Verbiest 2013), but allowing for the biases that affect low-frequency surveys, the measured distributions are closer to $\alpha = -1.8 \pm 0.2$ (where $S_\nu \propto \nu^\alpha$; Maron et al. 2000; Frail et al. 2016b; Kondratiev et al. 2016). Thus, there is a large frequency leverage arm for spectral index measurements when compared against existing centimetre surveys; those pulsars that may be weak or undetectable at 1.4 GHz are 50–1000× brighter at 150 MHz. Moreover, steep spectrum radio sources are rare. In Kimball & Ivezić (2008), fewer than 0.4 per cent of the radio sources have $\alpha < -1.8$ or less than 1 source per 140 deg². The only other known discrete radio class with similar spectral slopes is the luminous high-redshift galaxies, interesting in their own right but readily distinguished from pulsars for their kpc-sized extended structure resolved at arcsecond resolution (Miley & De Breuck 2008).

For this pilot study, we concentrated our search on the central 2:5 radius around Sgr A* where the GeV excess is strongest and where there are abundant ancillary data. We found a total of 220 sources in the TGSS ADR catalogue at 150 MHz above a threshold of 5σ . The brightest (faintest) source has a total flux of 35 Jy (43 mJy) and the median value of 250 mJy. As we are interested only in point sources and not extended sources (H II regions, supernova remnants, extragalactic sources, etc.), it would be standard practice to apply a cut-off based on the ratio of the total flux (S_t) to peak flux density (S_p) following Intema et al. (2017). However, this approach would likely eliminate real point sources due to the known enhanced scattering over large regions towards the galactic centre (e.g. Pynzar' & Shishov 2014). For example, a point source whose scattering diameter is 1 arcsec at 1 GHz would scale to be

44 arcsec at 150 MHz, larger than the 25 arcsec restoring beam of the TGSS ADR. We therefore used a more relaxed criterion empirically determined from the data as $S_t/S_p \leq 1.51$. Our final list consists of 166 point-like sources at 150 MHz.

As an initial pass at deriving spectral indices, we compared our source sample with the NRAO VLA Sky Survey (NVSS) catalogue at 1.4 GHz (Condon et al. 1998) using the TOPCAT software package (Taylor 2005). A total of 131 TGSS ADR sources had NVSS counterparts. For those remaining TGSS ADR sources without an NVSS counterpart we visually inspected the NVSS images in order to define a 3σ upper flux density limit based on the local noise properties. We reduced this initial candidate list further still by requiring that the initial two-point spectral index $\alpha < -1.4$. This condition was satisfied for 14 spectral index values and 5 spectral index limits, for a total of 19 sources.

Further refinements to the spectral index measurements of these 19 steep spectrum candidates were made by using SIMBAD to search for other imaging surveys of this region. Our primary catalogues were drawn from previous galactic surveys of compact radio sources and include Nord et al. (2004) made at 330 MHz with resolution, $\theta = 12$ arcsec \times 7 arcsec, and rms noise, $\sigma = 1.6$ mJy beam⁻¹, Zoonematkermani et al. (1990) at 1.5 GHz ($\theta = 5$ arcsec, $\sigma = 1$ – 2 mJy beam⁻¹), and Lazio & Cordes (1998a) at 1.5 GHz ($\theta = 5$ arcsec, $\sigma = 0.4$ mJy beam⁻¹) and 4.9 GHz ($\theta = 1.5$ arcsec, $\sigma = 0.4$ mJy beam⁻¹). In addition to our 19 TGSS ADR candidates, we carried out updated SIMBAD searches for additional flux density measurements on the five steepest spectrum sources of the 30 pulsar candidates identified in table 6 of Nord et al. (2004). New multifrequency spectral index measurements were derived from these added measurements. Many of the candidates had multiple flux density measurements at 1.4 GHz and flux density variations from one survey to the next were used to identify and eliminate likely resolved sources (although intrinsic variability could not be ruled out in some cases). Eight compact candidates remain with $\alpha < -1.4$. We chose the steepest spectrum sources with $\alpha < -1.7$ for follow-up. Only two radio sources within a 2:5 radius of Sgr A* satisfied this criterion: TGSS J174619.2–304010 and TGSS J175112.8–273723.

3 OBSERVATIONS AND DATA ANALYSIS

3.1 VLA

The two steep spectrum candidates identified in Section 2 were observed on 2016 July 8 with the Karl G. Jansky Very Large Array (VLA). The observations were carried out over a frequency range of 1–2 GHz with a standard set-up of sixteen 64 MHz subbands, with thirty-two 2 MHz channels in each subband (Project code = TDEM0009). The flux density and bandpass calibration used 3C 48 and phase calibration was carried out with the radio source J1751–2524. Data were calibrated using the NRAO pipeline and imaged with the CASA package. The VLA was in its B configuration giving an angular resolution in the images of approximately 5 arcsec.

At the higher angular resolution of the VLA our first candidate, TGSS J174619.2–304010, is resolved showing a head–tail morphology typical of extragalactic radio sources. Since we are searching for compact emission from pulsars we will not discuss this source any further. The VLA image for our second candidate, TGSS J175112.8–273723, is shown in Fig. 1. Two unresolved sources can be seen in this image but only one source is within the original beam of the GMRT detection at 150 MHz. A Gaussian fit to this source yields an improved J2000 position of RA = 17^h51^m12:65 and Dec. = $-27^\circ 37' 19''.8$ with an uncertainty of

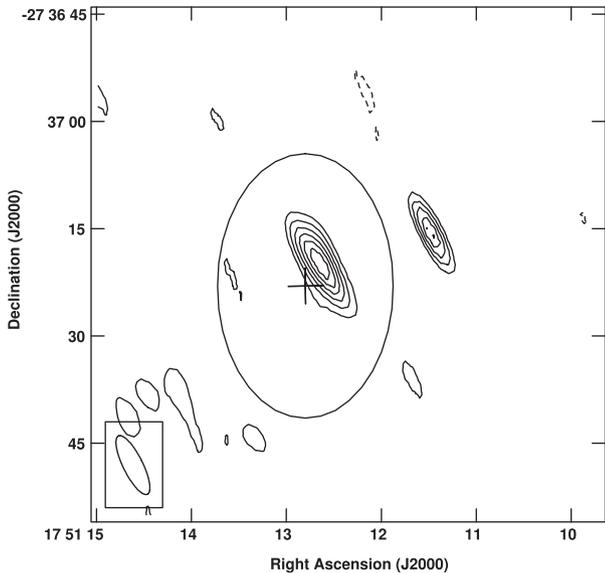


Figure 1. A continuum VLA image at 1.5 GHz of two radio sources. The steep spectrum pulsar candidate lies near the centre with an ellipse showing the original TGSS ADR beam and a cross indicating the estimate 1σ error in this position. A second background radio source lies to the northeast. The VLA synthesized beam size is shown in the lower left corner. The contours are in units of the rms noise of $50 \mu\text{Jy beam}^{-1}$ starting at $-3, 3, 5, 7, 9, 11, 13$ and 15 .

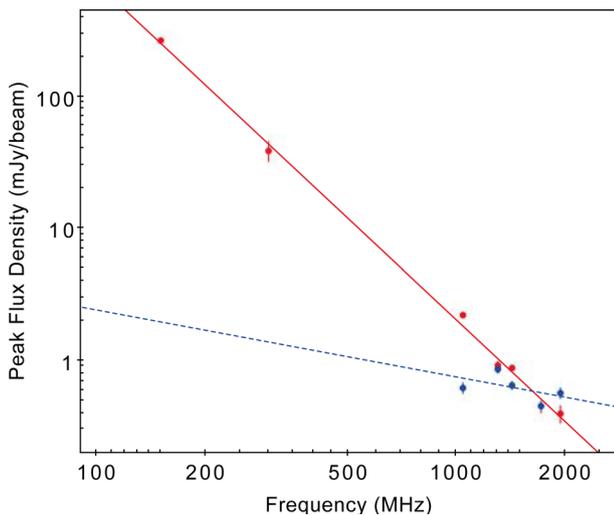


Figure 2. Continuum radio spectra of the two VLA point sources, with the steep spectrum pulsar candidate in red and the fainter, flat spectrum source in blue. A least-squares fit to the data gives power-law slopes of $\alpha = -0.52 \pm 0.35$ (blue dashed) and $\alpha = -2.55 \pm 0.08$ (red solid).

± 0.3 arcsec. We averaged the visibility data into five separate sub-bands and imaged each, measuring the peak and integrated flux density for both sources. The resulting spectra are also shown in Fig. 2. A power-law fit to the peak flux values gives $\alpha = -0.52 \pm 0.35$ and $\alpha = -2.55 \pm 0.08$ for the background source and the pulsar candidate, respectively. Fitting to the total flux density instead of the peak flux gives similar slopes. The TGSS source has a spectral slope that lies on the extreme tail of the pulsar spectral index distribution (Maron et al. 2000; Bates et al. 2013), while the slope of the other source is more typical of extragalactic sources.

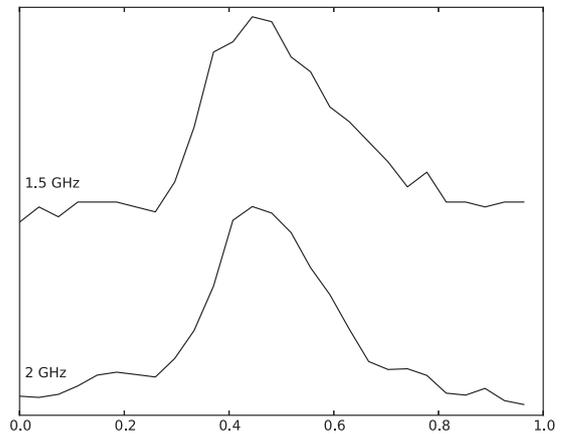


Figure 3. Top: normalized average pulse profile of PSR J1751–2737 at 1.5 GHz from our single *L*-band GBT observation. Bottom: pulse profile at 2.0 GHz obtained by averaging the normalized pulse profiles of our three *S*-band GBT observations.

3.2 GBT

We observed TGSS J175112.8–273723 with the GBT (Project code 16B-384) at three epochs and three frequencies, 1.5 GHz (*L* band), 2.0 GHz (*S* band) and 5.0 GHz (*C* band). All observations used the GUPPI backend, and Table 2 lists the details of each. While the GUPPI bandwidth was 800 MHz for all observations in the table, the 2.0 GHz GBT receiver has permanent filters that reduce its effective bandwidth to 610 MHz.

We searched the 2016 October 4 *S*-band observation for pulsed signals using the PRESTO software¹ and a list of 4770 trial dispersion measures (DMs) in the range $0\text{--}2500 \text{ pc cm}^{-3}$. A pulsed signal with a period of 2.23 ms was detected at trial $\text{DM} = 260.58 \text{ pc cm}^{-3}$. We identified it as astrophysical based on its wide-band signature and the fact that it exhibits a peak in signal to noise versus trial DM characteristic of dispersed pulsars (see e.g. figs 3 and 4 in Cordes & McLaughlin 2003). The period and DM were verified in subsequent *S*-band and *L*-band observations; our *C*-band observation did not yield a detection. The new pulsar is angularly close to the Galactic Centre with galactic coordinates $l, b = 1.76, -0.38$, but based on its DM and the NE2001 model of ionized gas in the Galaxy (Cordes & Lazio 2002), it is a foreground object at a distance of ~ 4 kpc. The newer YMW16 model (Yao, Manchester & Wang 2017) gives a distance of ~ 3.4 kpc. Table 3 lists the pulsar parameters.

We processed our GBT observations with `rfifind` and `prepfold` from PRESTO to excise radio frequency interference (RFI) and produce average pulsar profiles. Fig. 3 shows the normalized pulse profile (i.e. scaled to a peak of unity) from the single *L*-band observation and the result from averaging the normalized profiles of the three *S*-band observations. The FWHM pulse widths are 0.56 and 0.52 ms at 1.5 and 2.0 GHz, respectively.

We fit the two normalized average profiles with a Gaussian convolved with an exponential with a decay time corresponding to the scattering broadening time, τ_s . The *S*-band best-fitting Gaussian $\sigma = 0.18$ ms and $\tau_s = 0.19$ ms, and the *L*-band best-fitting $\sigma = 0.15$ ms and $\tau_s = 0.30$ ms. For comparison, we also computed the expected τ_s for the two observing frequencies and the pulsar DM using the NE2001 model of ionized gas in the Galaxy.² This yielded τ_s values of 0.07 and 0.26 ms for *S* band and *L* band,

¹ <http://www.cv.nrao.edu/~sransom/presto/>

² <https://www.nrl.navy.mil/rsd/RORF/ne2001>

Table 1. Measured flux values for TGSS J175112.8–273723. Columns represent centre observing frequency, peak flux and total flux density.

Frequency (MHz)	S_p (mJy beam $^{-1}$)	S_t (mJy)
150	268.4 ± 30.8	260.8 ± 37.6
330	38.8 ± 6.5	48.8 ± 9.8
1040	2.218 ± 0.066	2.170 ± 0.150
1296	0.928 ± 0.064	0.871 ± 0.147
1424	0.889 ± 0.045	0.801 ± 0.101
1712	0.453 ± 0.044	0.625 ± 0.107
1936	0.393 ± 0.054	0.379 ± 0.151

respectively. Our results are within the uncertainties of models estimating τ_s based on measured pulsar DM (Cordes & Lazio 2003, Bhat et al. 2004).

We estimate the flux density of PSR J1751–2737 based on the radiometer equation, the average pulse profile of each detection, the parameters in Table 2 and the background sky temperature T_{sky} at each frequency. In order to find T_{sky} at each of our observing frequencies, we scale from the observed value of 3.1 K at 2695 MHz (Reich et al. 1990) to each frequency using a power law with an index of -2.6 (Reich & Reich 1988).³ We obtain $T_{\text{sky}, 1.5\text{GHz}} = 14$ K, $T_{\text{sky}, 2\text{GHz}} = 7$ K and $T_{\text{sky}, 5\text{GHz}} = 0.6$ K.

From our 1.5 GHz observation we obtain a period-averaged flux density of 0.32 mJy. However, the observation was plagued by strong and pervasive RFI that was difficult to excise without masking most of the data. This flux density value should be taken as a lower limit. From our 2.0 GHz observations we obtain period-averaged flux densities of 0.35, 0.16 and 0.38 mJy. While the S-band receiver bandwidth is significantly cleaner than the L-band bandwidth, the outlying middle measurement may be due to imperfectly cleaned RFI. Regular observations during our planned timing campaign of the new millisecond pulsar (MSP) will be able to distinguish between this possibility and propagation effects such as refractive scintillation. The average flux density of the three 2.0 GHz observations, 0.30 mJy, agrees very well with the extrapolation of the VLA measurements in Table 1 and Fig. 2. From our non-detection at 5.0 GHz, we calculate an upper limit of 0.04 mJy, assuming a detection threshold of 10σ . From the C-band upper limit and the average S-band flux density, we calculate an upper limit on the spectral index of the pulsar, $\alpha < -2.2$, again in good agreement value derived from interferometric data.

4 DISCUSSION AND CONCLUSIONS

PSR J1751–2737 appears to be an isolated recycled pulsar located within the disc of our Galaxy. A binary pulsar would have been expected to show some change in its period over the three-week interval so that it was observed with the GBT, while none was observed (Table 3). However, until longer term timing is conducted, a long-period binary or one with a low mass fraction cannot be ruled out. The pulsar’s radio luminosity at 1.4 GHz, defined in the usual way as $L_{1.4} = S_t \times d^2$, and using the NE2001 distance is 13 mJy kpc². This value of $L_{1.4}$ is on the high end of the lognormal luminosity function for MSPs in globular clusters (Bagchi, Lorimer

Table 2. GBT observations of PSR J1751–2737. Columns represent centre observing frequency, effective bandwidth (the smaller of the backend bandwidth and unfiltered receiver bandwidth), sampling time, gain, system temperature, integration time and the presence of a pulsed signal detection. T_{sys} includes instrumental contributions as well as the cosmic microwave background and atmospheric emission.

Date (2016)	f_c (GHz)	$\Delta\nu$ (MHz)	Δt (μ s)	G (K Jy $^{-1}$)	T_{sys} (K)	T_{obs} (h)	Detect?
October 4	2.0	610	81.92	2.0	20	0.79	Y
October 11	2.0	610	81.92	2.0	20	1.09	Y
October 11	5.0	800	81.92	1.9	18	2.89	N
October 23	1.5	800	81.92	1.9	20	0.96	Y
October 23	2.0	610	81.92	2.0	20	1.00	Y

& Chennamangalam 2011) and is a bright outlier in the existing sample of isolated MSPs in the disc (Burgay et al. 2013). The dispersion measure distance and the limits on temporal scattering (Section 3.2) favour an origin in the galactic plane. For the pulsar to be located in the bulge at ~ 8.5 kpc the NE2001 model predicts a DM = 655 pc cm $^{-3}$ and τ_s at 2 GHz of 2.4 ms, values much higher than we observe. Thus, PSR J1751–2737 is not a member of the putative bulge population that is thought to be responsible for the Fermi excess. There are no discrete gamma-ray sources in this direction either in the 3FGL catalogue or in catalogues of the galactic centre region (Acero et al. 2015; Ajello et al. 2016). A search for gamma-ray pulsations may still be worthwhile since the complex background models in this direction make it difficult to robustly identify discrete sources.

Most past sensitive centimetre radio pulsar surveys did not search a region large enough to detect PSR J1751–2737. Owing to the large amount of telescope time required for high-frequency single-dish surveys, some searches have concentrated within a 1° radius centred on SgrA* (Johnston et al. 2006; Deneva et al. 2009), while others have focused deep integrations within the central few parsecs around SgrA* (Macquart et al. 2010; Siemion et al. 2013; Eatough et al. 2013). The lone exception is the on-going High Time Resolution Universe (HTRU) survey of the southern sky with the Parkes multibeam receiver at 1.4 GHz, which has covered the significant parts of the Fermi excess area. In particular, the low-latitude component of the southern HTRU survey is making hour-long integrations with the goal of searching the galactic plane in latitude $|b| < 3.5^\circ$ and longitude $30^\circ < l < 280^\circ$ (Keith et al. 2010). HTRU has searched the direction towards PSR J1751–2737 but the pulsar does not appear in recent lists of detections (Ng et al. 2015). This is somewhat surprising given the experimental parameters of the survey and the similarity between the GBT and Parkes sensitivities (see Table 2 and Keith et al. 2010). A non-detection above threshold could be explained if PSR J1751–2737 was beyond the half-power point of a single feed, especially if it was located near an outer feed where the telescope gain is lower. In any case, with the position and period from Table 3 a renewed HTRU archival search would be worthwhile. When compared to the GBT data, earlier HTRU measurements could be used to help constrain \dot{E} and other parameters.

While PSR J1751–2737 is not a bulge MSP, this pilot project has shown that interferometric imaging observations have a role in constraining the contribution of young and recycled pulsars to the gamma-ray excess around the galactic centre. Calore et al. (2016) have looked at the prospects for detecting a bulge population of MSPs in some detail. They show that while existing deep surveys such as the HTRU are not well suited to bulge detections, future

³ PSR J1751–2737 is outside the Reich & Reich (1988) spectral index map by $\sim 7^\circ$; however, the spectral index of T_{sky} does not vary significantly on that spatial scale in the vicinity of the pulsar.

Table 3. Properties of PSR J1751–2737: coordinates are from VLA imaging data, and all other values are derived from GBT time domain data.

Parameter	Value
Name	J1751–2737
Right ascension (J2000)	17 ^h 51 ^m 12 ^s .65(2)
Declination (J2000)	−27°37′19″.8(3)
Rotation period P (ms)	2.23
Dispersion measure DM (pc cm ^{−3})	260
NE2001 distance D (kpc)	4.0
YMW16 distance D (kpc)	3.4
$S_{1.5\text{ GHz}}$ (mJy)	>0.32
$S_{2\text{ GHz}}$ (mJy)	0.30
$S_{5\text{ GHz}}$ (mJy)	<0.04
Spectral index α ($S \propto \nu^\alpha$)	<−2.2
$\sigma_{1.5\text{ GHz}}$ (ms)	0.15
$\tau_{s, 1.5\text{ GHz}}$ (ms)	0.30
$\sigma_{2\text{ GHz}}$ (ms)	0.18
$\tau_{s, 2\text{ GHz}}$ (ms)	0.19

large area radio pulsation searches have the potential to yield dozens of detections. The disadvantage of this direct approach, however, is that it is currently prohibitively expensive, either in telescope time or in computational resources. Here we used a hybrid approach. We undertook interferometric imaging at low frequencies where a single 15-min integration at 150 MHz covered a field of view of 5 deg². Pulsar candidates were identified as steep spectrum point sources. This approach is sensitive to only phase-averaged flux density. One limitation of this method is that it is not sensitive to flat spectrum pulsars. In the current pulsar catalogue, one third of all pulsars have a spectral index $\alpha > -1.6$. This disadvantage is balanced against other advantages. The method makes no assumptions about the pulsar period, the dispersion measure, temporal scattering or binarity. These quantities were searched over from a time series taken in a deep single-pointing pulsation search carried out at higher frequencies.

We can make some rough comparisons between a direct pulsation search and the hybrid technique. Calore et al. (2016) describe a pulsation search with a 100-m class single-dish radio telescope at 1.4 GHz. Their region with the highest yield is 5° above (or below) the Galactic Centre. The surface density of radio bright bulge MSPs (defined as ≥ 10 μJy at 1.4 GHz) is still large (4.7 ± 1.5 deg^{−2}) but the sky brightness temperature, the scattering and dispersive smearing are all sharply reduced compared to the Galactic Centre. In a 250-h experiment they show that a 100-m telescope can carry out a sensitive survey of a 2° × 2° area, detecting 1.7 bulge MSPs or 0.43 detections deg^{−2}. A hybrid approach would begin with a deep interferometric image of the same region. For example, a single 1-h pointing of the upgraded GMRT 325 MHz could achieve a thermal noise⁴ (5σ) of 15 $\mu\text{Jy beam}^{-1}$ over a field of view of 1.4 deg². Pulsars are 10 times brighter at 325 MHz than at 1.4 GHz, so this experiment would detect *all* radio bright bulge MSPs. Since approximately two thirds of these would be selected based on their steep spectrum, imaging over the GMRT field of view results in 4.5 MSP candidates or 3.1 detections deg^{−2}. (Note that the smaller dishes of the VLA would require more integration time to reach similar noise levels but they would image a field of view three times larger.) Follow-up single-dish pulsation searches would require several hours to confirm pulsations from each of these 4.5 candidates but they would be made at frequencies higher than 1.4 GHz, where

the scattering and dispersive effects are reduced by λ^4 and λ^2 , respectively. These numbers are only approximate but they show that the hybrid technique might be used to get an order of magnitude yield of MSP bulge detections, with a concomitant reduction in the total observing time.

This hybrid approach could be used in the future to efficiently search for a bulge population of pulsars. Imaging searches for compact steep spectrum source towards the galactic centre region have been carried out in the past (Lazio & Cordes 1998a; Nord et al. 2004), but a new generation of low-temperature, broad-band feeds motivates a re-thinking of new surveys. Instantaneous, wide bandwidth measurements of flux densities, while subject to some uncertainties (Rau, Bhatnagar & Owen 2016), are preferable to measuring the spectral index by comparing images from different telescopes with large differences in angular resolution and different calibration schemes, and often taken at different times so that variability can produce false detections. For example, many of the initial steep spectrum candidates identified in Nord et al. (2004) turn out to be false positives due to resolution effects or variability. TGSS J175112.8–273723 on the other hand could have been identified as a strong pulsar candidate based solely on the follow-up VLA measurements taken from 1 to 2 GHz (Table 1).

The VLA L -band system (1–2 GHz) and the upgraded P -band system at the GMRT (250–500 MHz) have the large fractional bandwidths ($\delta\nu/\nu = 33$ per cent) required to detect steep spectrum candidates in band. They also have the wide instantaneous field of view needed to efficiently image the entire region of the *Fermi* excess with arcsecond resolution. Our choice of 150 MHz for the pilot study was suboptimal since the bandwidth is narrow ($\delta\nu/\nu = 10$ per cent) while the temperature of the diffuse sky background is large and rises as $\nu^{-2.6}$, i.e. steeper than most pulsars. There is also free-free absorption from the patchy, ionized gas at the galactic centre that attenuates pulsar emission at this frequency (Pynzar’ & Shishov 2014). With the proper choice of channel widths and integration times for the interferometer set-up, pulsar candidates could be further identified based on their diffractive scintillations as measured in variance images (Dai et al. 2016). In the mean time, before a full hybrid survey is carried out, we agree with O’Leary et al. (2016) and Calore et al. (2016) that a search should be made for bulge pulsar candidates towards the ‘hotspots’ identified by Bartels et al. (2016) and Lee et al. (2016). The detection of PSR J1751–2737 shows that deep imaging-based continuum with targeted pulsation searches offers a complementary approach to existing blind pulsation surveys.

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⁴ N. Kanekar, private communication

REFERENCES

- Abazajian K. N., 2011, *J. Cosmol. Astropart. Phys.*, 3, 010
- Acerio F. et al., 2015, *ApJS*, 218, 23
- Ajello M. et al., 2016, *ApJ*, 819, 44
- Bagchi M., Lorimer D. R., Chennamangalam J., 2011, *MNRAS*, 418, 477
- Bartels R., Krishnamurthy S., Weniger C., 2016, *Phys. Rev. Lett.*, 116, 051102
- Bates S. D., Lorimer D. R., Verbiest J. P. W., 2013, *MNRAS*, 431, 1352
- Bhat N. D. R., Cordes J. M., Camilo F., Nice D. J., Lorimer D. R., 2004, *ApJ*, 605, 759
- Brandt T. D., Kocsis B., 2015, *ApJ*, 812, 15
- Broderick J. W. et al., 2016, *MNRAS*, 459, 2681
- Burgay M. et al., 2013, *MNRAS*, 433, 259
- Calore F., Cholis I., Weniger C., 2015, *J. Cosmol. Astropart. Phys.*, 3, 038
- Calore F., Di Mauro M., Donato F., Hessels J. W. T., Weniger C., 2016, *ApJ*, 827, 143
- Condon J. J., Cotton W. D., Greisen E. W., Yin Q. F., Perley R. A., Taylor G. B., Broderick J. J., 1998, *AJ*, 115, 1693
- Cordes J. M., Lazio T. J. W., 1997, *ApJ*, 475, 557
- Cordes J. M., Lazio T. J. W., 2002, preprint ([astro-ph/0207156](https://arxiv.org/abs/astro-ph/0207156))
- Cordes J. M., Lazio T. J. W., 2003, preprint ([astro-ph/0301598](https://arxiv.org/abs/astro-ph/0301598))
- Cordes J. M., McLaughlin M. A., 2003, *ApJ*, 596, 1142
- Dai S., Johnston S., Bell M. E., Coles W. A., Hobbs G., Ekers R. D., Lenc E., 2016, *MNRAS*, 462, 3115
- Daylan T., Finkbeiner D. P., Hooper D., Linden T., Portillo S. K. N., Rodd N. L., Slatyer T. R., 2016, *Physics of the Dark Universe*, 12, 1
- Deneva J. S., Cordes J. M., Lazio T. J. W., 2009, *ApJ*, 702, L177
- Eatough R. P. et al., 2013, *Nature*, 501, 391
- Eatough R. P., Kramer M., Klein B., Karuppusamy R., Champion D. J., Freire P. C. C., Wex N., Liu K., 2013, in van Leeuwen J., ed., *Proc. IAU Symp. 291, Neutron Stars and Pulsars: Challenges and Opportunities after 80 Years*. Cambridge Univ. Press, Cambridge, 382
- Frail D. A., Diamond P. J., Cordes J. M., van Langevelde H. J., 1994, *ApJ*, 427, L43
- Frail D. A., Mooley K. P., Jagannathan P., Intema H. T., 2016a, *MNRAS*, 461, 1062
- Frail D. A., Jagannathan P., Mooley K. P., Intema H. T., 2016b, *ApJ*, 829, 119
- Hooper D., Goodenough L., 2011, *Phys. Lett. B*, 697, 412
- Intema H. T., Jagannathan P., Mooley K. P., Frail D. A., 2017, *A&A*, 598, A78
- Johnston S., Kramer M., Lorimer D. R., Lyne A. G., McLaughlin M., Klein B., Manchester R. N., 2006, *MNRAS*, 373, L6
- Keith M. J. et al., 2010, *MNRAS*, 409, 619
- Kimball A. E., Ivezić Ž., 2008, *AJ*, 136, 684
- Kondratiev V. I. et al., 2016, *A&A*, 585, A128
- Lazio T. J. W., Cordes J. M., 1998a, *ApJS*, 118, 201
- Lazio T. J. W., Cordes J. M., 1998b, *ApJ*, 505, 715
- Lazio T. J. W., Cordes J. M., 2008, *ApJS*, 174, 481
- Lee S. K., Lisanti M., Safdi B. R., Slatyer T. R., Xue W., 2016, *Phys. Rev. Lett.*, 116, 051103
- Macquart J.-P., Kanekar N., 2015, *ApJ*, 805, 172
- Macquart J.-P., Kanekar N., Frail D. A., Ransom S. M., 2010, *ApJ*, 715, 939
- Maron O., Kijak J., Kramer M., Wielebinski R., 2000, *A&A*, 147, 195
- Miley G., De Breuck C., 2008, *A&AR*, 5, 67
- Mirabal N., 2013, *MNRAS*, 436, 2461
- Mori K. et al., 2013, *ApJ*, 770, L23
- Muno M. P., Pfahl E., Baganoff F. K., Brandt W. N., Ghez A., Lu J., Morris M. R., 2005, *ApJ*, 622, L113
- Navarro J., de Bruyn A. G., Frail D. A., Kulkarni S. R., Lyne A. G., 1995, *ApJ*, 455, L55
- Ng C. et al., 2015, *MNRAS*, 450, 2922
- Nord M. E., Lazio T. J. W., Kassim N. E., Hyman S. D., LaRosa T. N., Brogan C. L., Duric N., 2004, *AJ*, 128, 1646
- O’Leary R. M., Kistler M. D., Kerr M., Dexter J., 2016, *Phys. Rev. D*, preprint ([arXiv:1601.05797](https://arxiv.org/abs/1601.05797))
- Pfahl E., Loeb A., 2004, *ApJ*, 615, 253
- Pynzar’ A. V., Shishov V. I., 2014, *Astron. Rep.*, 58, 427
- Rau U., Bhatnagar S., Owen F. N., 2016, *AJ*, 152, 124
- Reich P., Reich W., 1988, *A&AS*, 74, 7
- Reich W., Fürst E., Reich P., Reif K., 1990, *A&AS*, 85, 663
- Shannon R. M., Johnston S., 2013, *MNRAS*, 435, L29
- Siemion A. et al., 2013, in van Leeuwen J., ed., *Proc. IAU Symp. 291, Neutron Stars and Pulsars: Challenges and Opportunities after 80 Years*. Cambridge Univ. Press, Cambridge, p. 57
- Taylor M. B., 2005, in Shopbell P., Britton M., Ebert R., eds, *ASP Conf. Ser. Vol. 347, Astronomical Data Analysis Software and Systems XIV*. Astron. Soc. Pac., San Francisco, p. 29
- van Eldik C., 2015, *Astropart. Phys.*, 71, 45
- Vitale V., Morselli A., for the Fermi/LAT Collaboration 2009, preprint ([arXiv:0912.3828](https://arxiv.org/abs/0912.3828))
- Wang Q. D., Lu F. J., Gotthelf E. V., 2006, *MNRAS*, 367, 937
- Weniger C., 2012, *J. Cosmol. Astropart. Phys.*, 8, 007
- Wharton R. S., Chatterjee S., Cordes J. M., Deneva J. S., Lazio T. J. W., 2012, *ApJ*, 753, 108
- Yao J. M., Manchester R. N., Wang N., 2017, *ApJ*, 835, 29
- Yuan Q., Zhang B., 2014, *J. High Energy Phys.*, 3, 1
- Zoonematkermani S., Helfand D. J., Becker R. H., White R. L., Perley R. A., 1990, *ApJS*, 74, 181

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